Dark matter crisis on small scales?

Carlos S. Frenk
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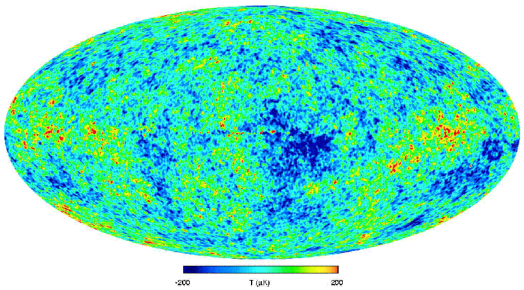
Durham

Four problems for CDM on small scales?

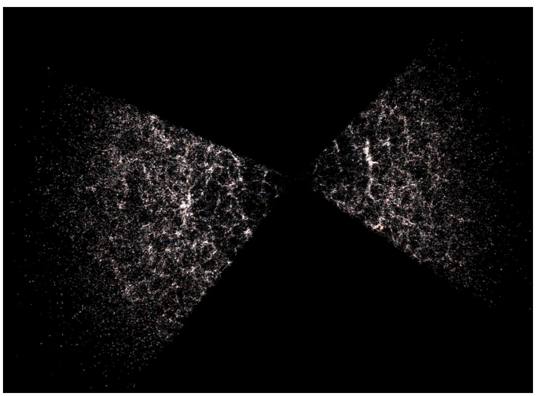
1. The “missing satellites” problem
2. The “too-big-to-fail” problem
3. The “core-cusp” problem
4. The “satellite disk” problem

University of Durham

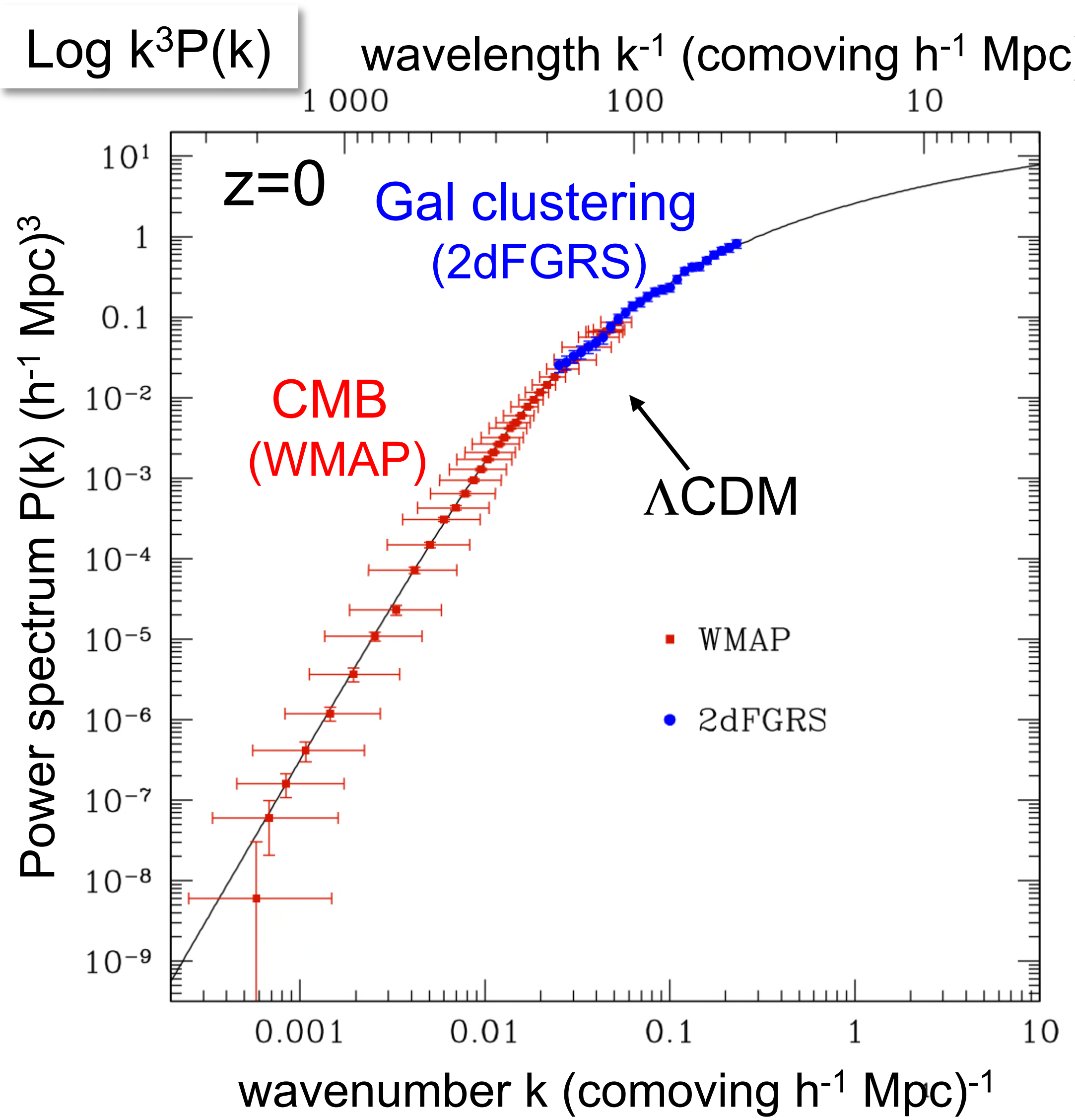
The cosmic power spectrum: from the CMB to the 2dFGRS



z~1000



z~0



⇒ ΛCDM provides an excellent description of mass power spectrum from 10-1000 Mpc

Sanchez et al 06

The cosmic power spectrum: from the CMB to the 2dFGRS

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Free streaming →

$\lambda_{cut} \propto m_x^{-1}$

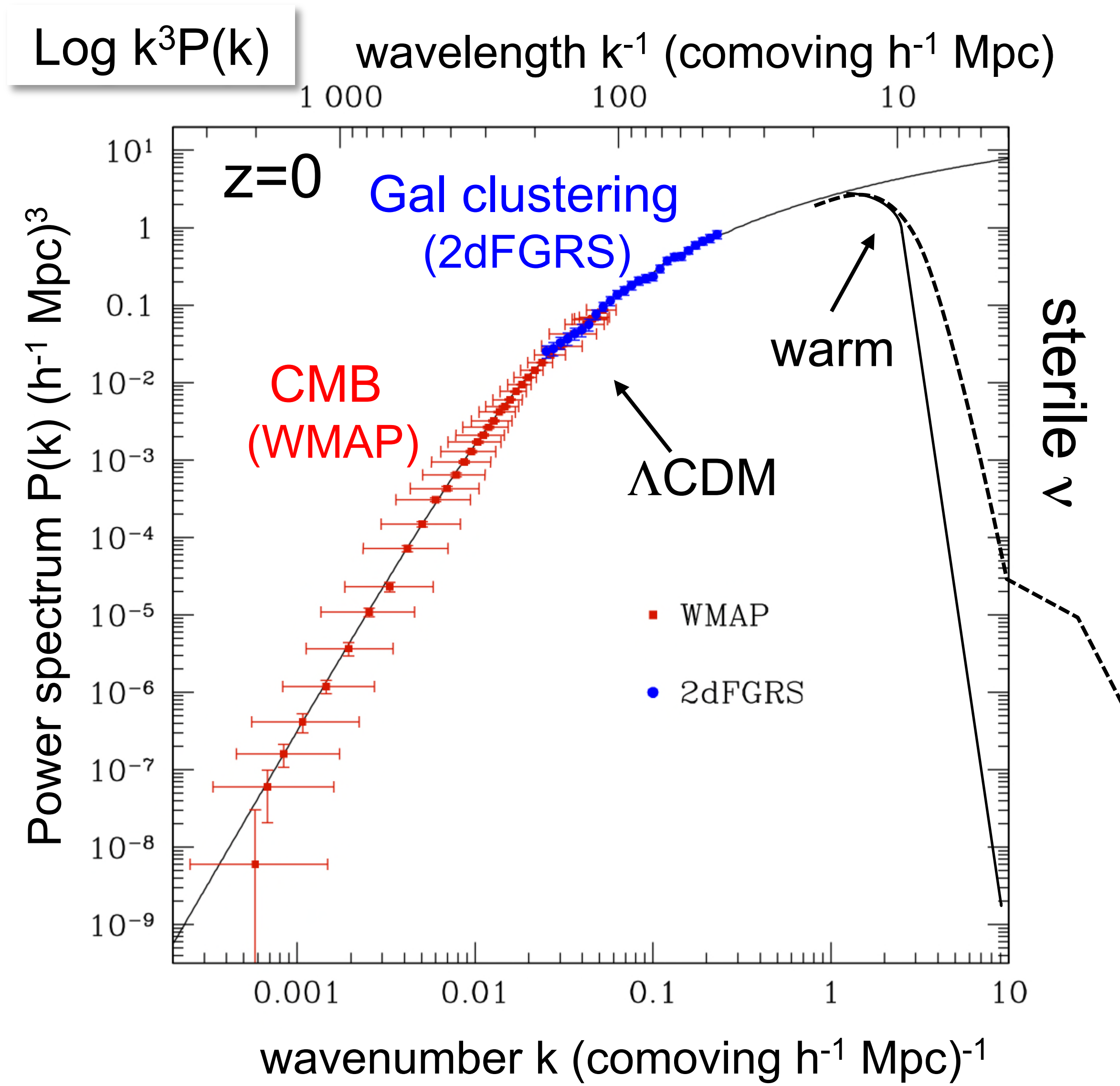
for thermal relic

$m_{CDM} \sim 100GeV$

susy; $M_{cut} \sim 10^{-6}\, M_o$

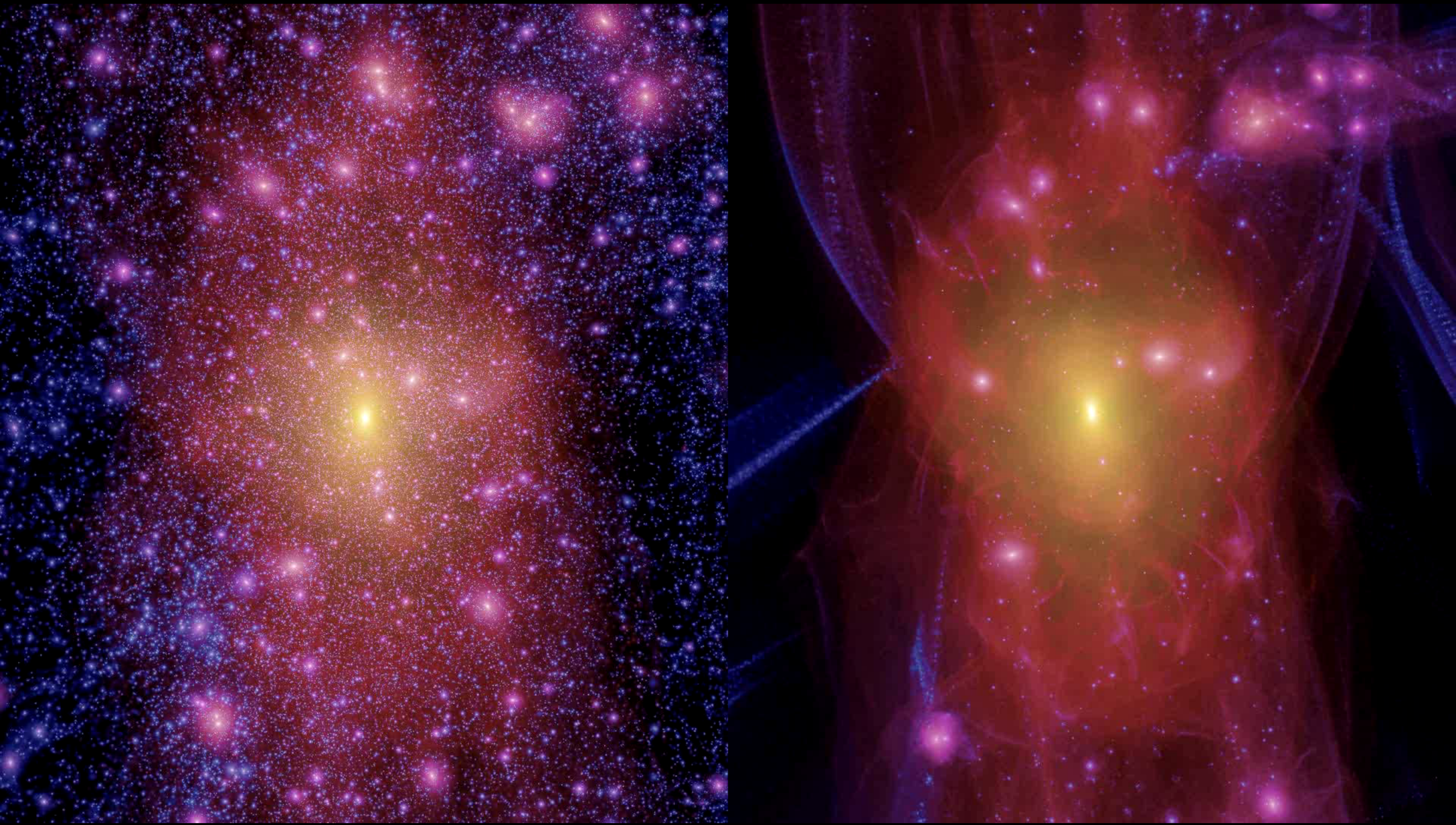
$m_{WDM} \sim$ few keV

sterile ν; $M_{cut} \sim 10^{9}\, M_o$



cold dark matter

warm dark matter



Lovell, Eke, Frenk, Gao, Jenkins, Wang, White, Theuns, Boyarski & Ruchayskiy ‘12

Four problems for CDM on small scales?

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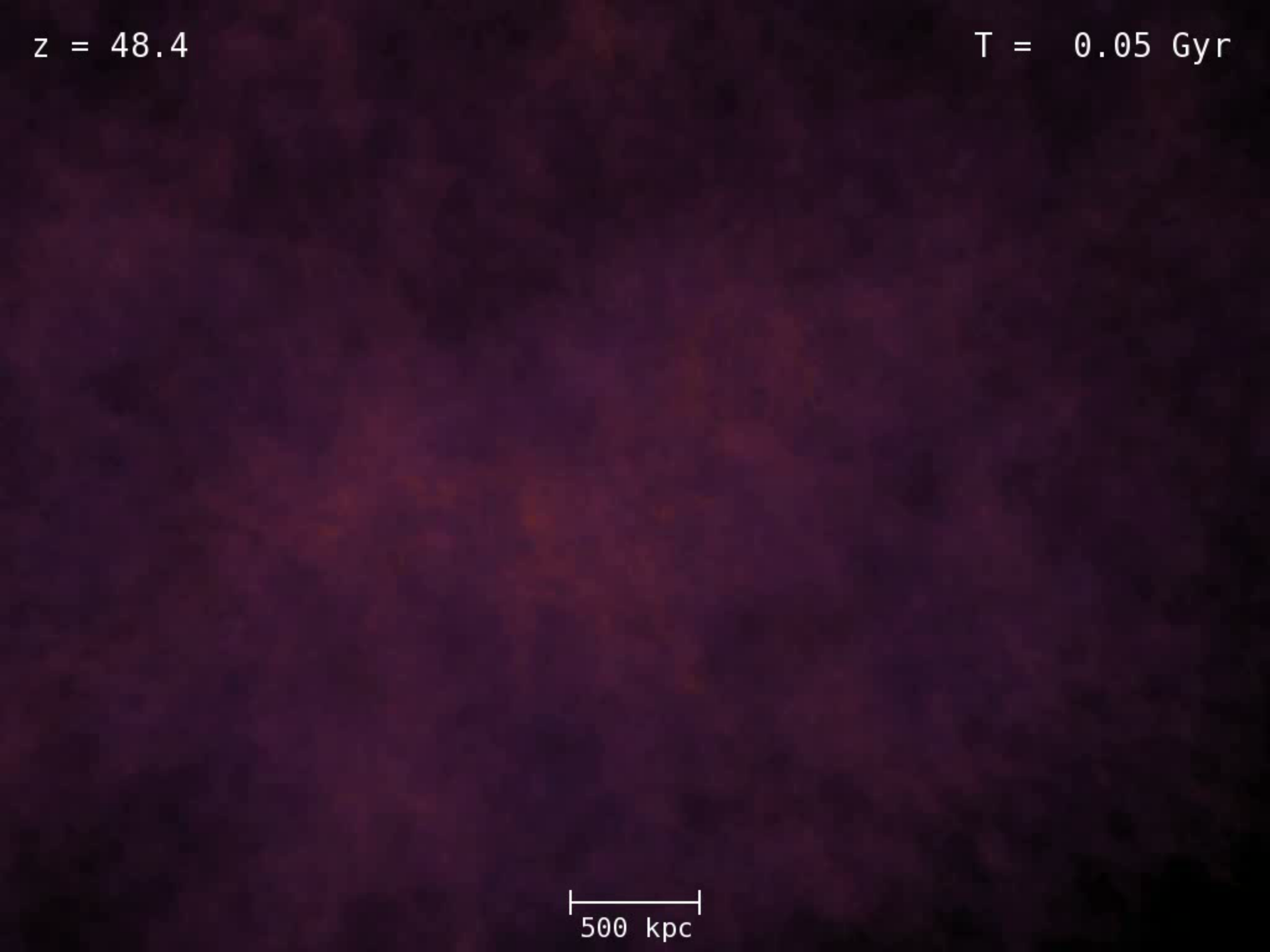
These problems have all been identified in N-body simulations that follow only dark matter

Need to consider “baryon effects”


z = 48.4
T = 0.05 Gyr
500 kpc

VIRGO

icc.dur.ac.uk/Eagle

“Evolution and assembly of galaxies and their environment”

THE EAGLE PROJECT

Durham: Richard Bower, Michelle Furlong, Carlos Frenk, Matthieu Schaller, James Trayford, Yelti Rosas-Guevara, Tom Theuns, Yan Qu, John Helly, Adrian Jenkins.

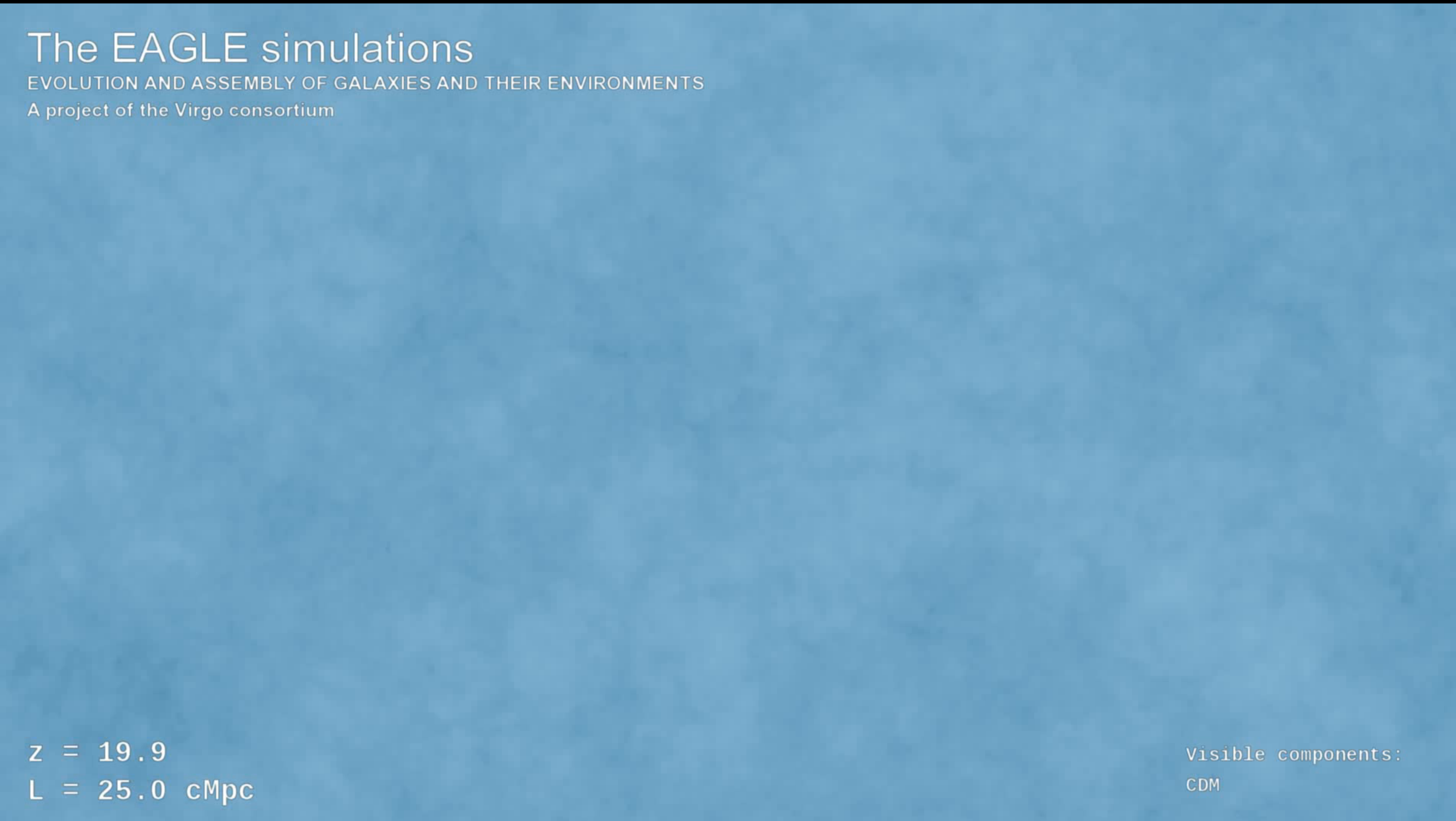
Leiden: Rob Crain, Joop Schaye.

Other: Claudio Dalla Vecchia, Ian McCarthy, Craig Booth…

+ Virgo Consortium

Durham University VIRGO ICC DiRAC PRACE

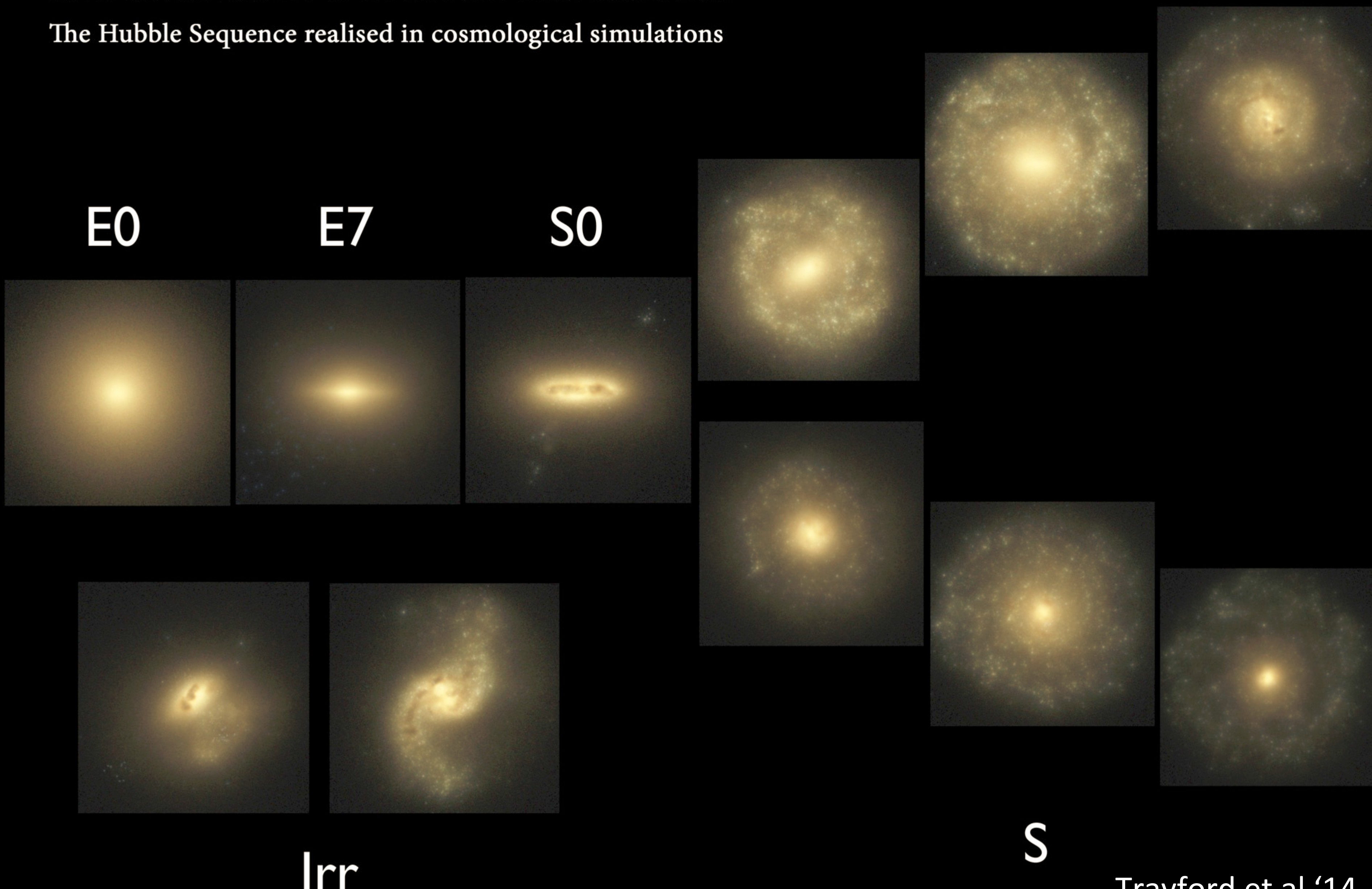
The EAGLE simulation is a project by the Virgo Consortium.
Images courtesy of Richard Bower.


The EAGLE simulations
EVOLUTION AND ASSEMBLY OF GALAXIES AND THEIR ENVIRONMENTS
A project of the Virgo consortium
z = 19.9
L = 25.0 cMpc
Visible components:
CDM

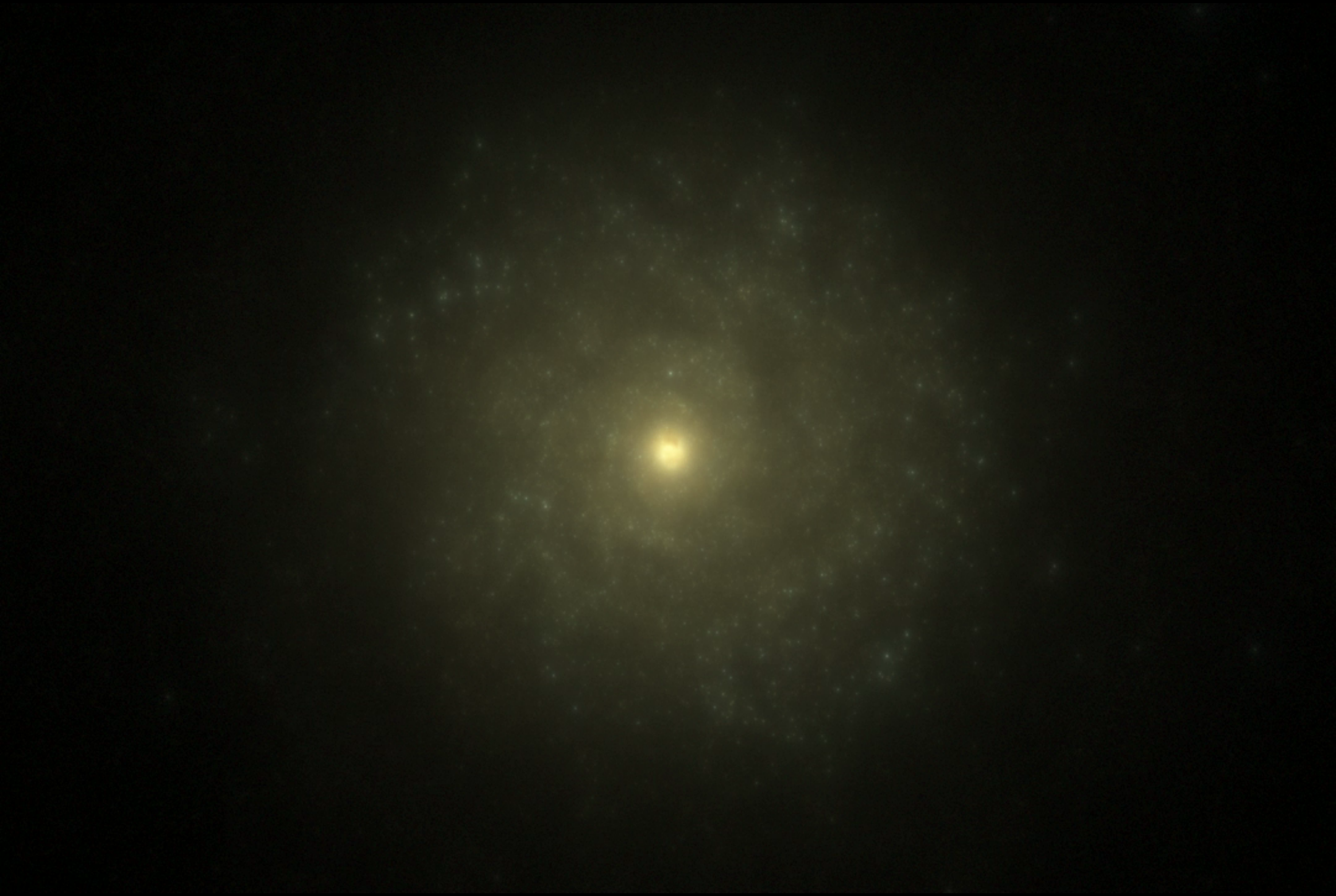
The Eagle Simulations

EVOLUTION AND ASSEMBLY OF GALAXIES AND THEIR ENVIRONMENTS

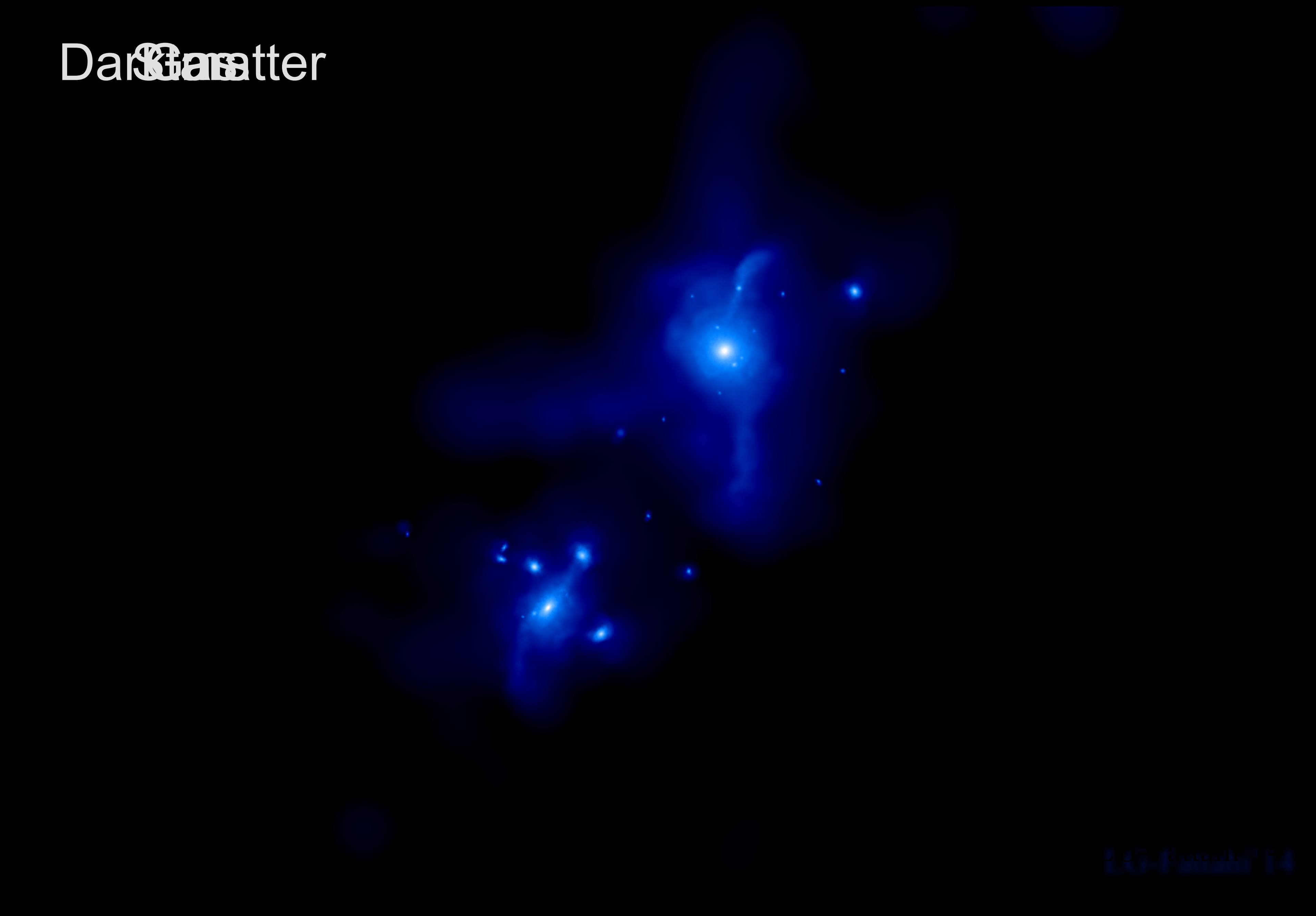
The Hubble Sequence realised in cosmological simulations

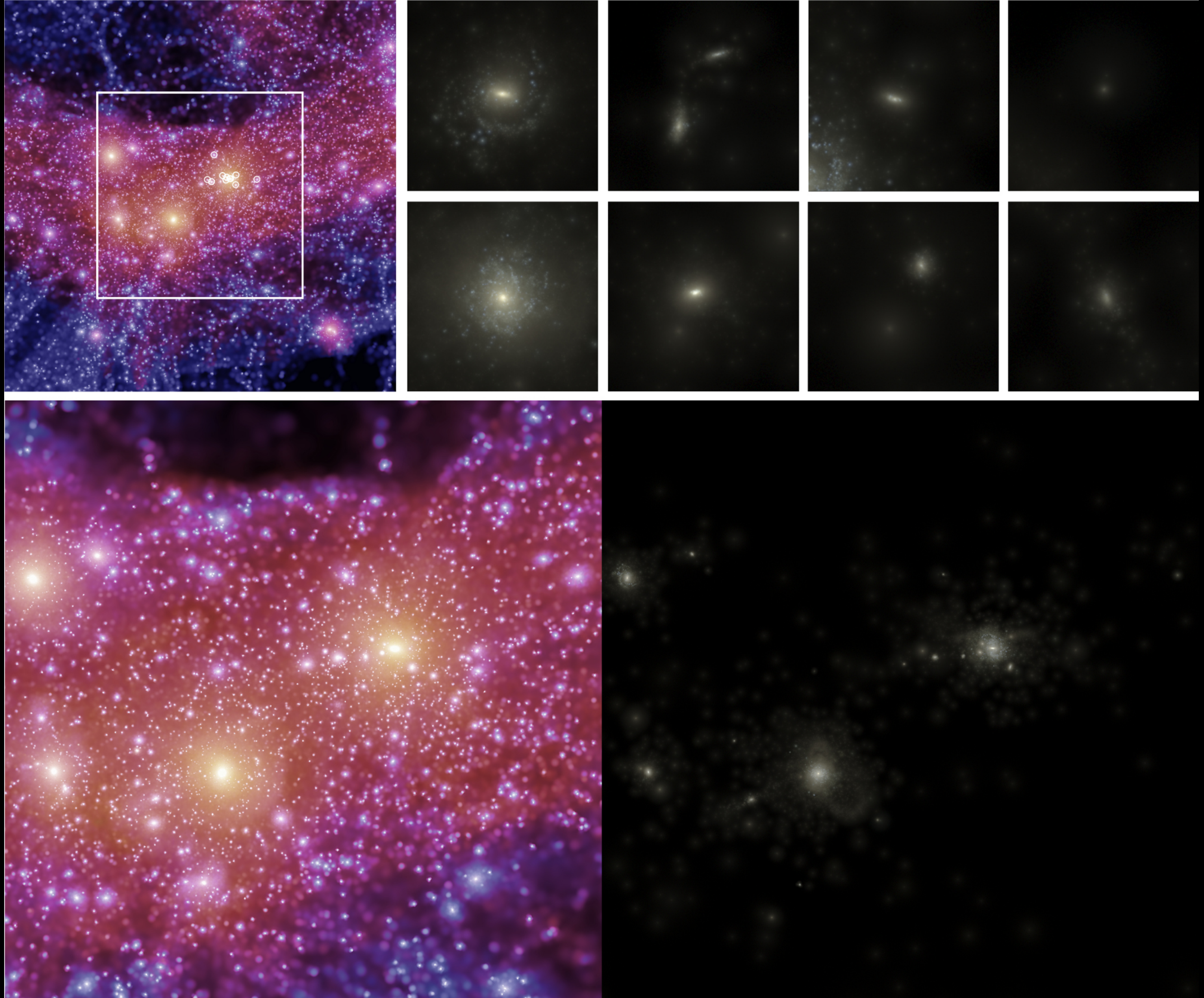


Trayford et al '14



EAGLE full hydro Local Group simulations





Four problems for CDM on small scales?

1. The “missing satellites” problem
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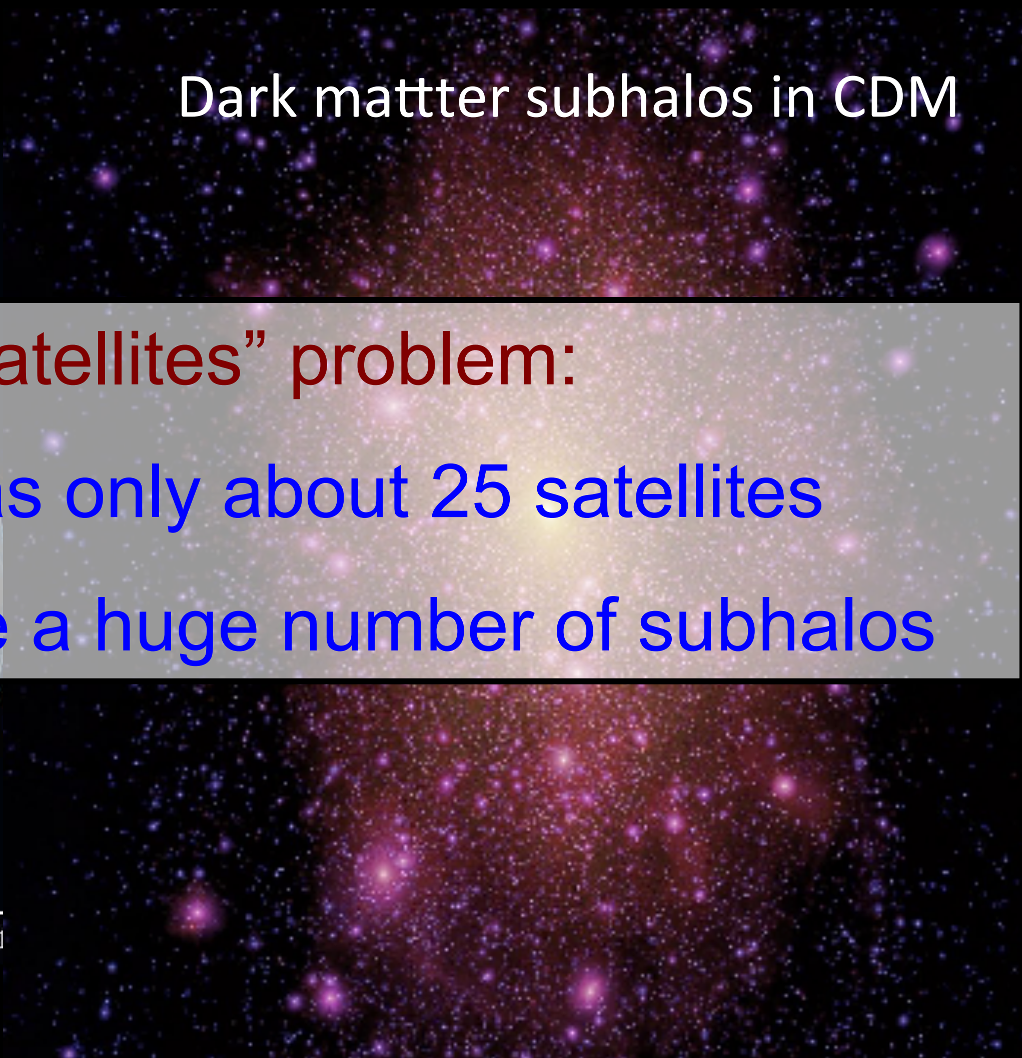
The “missing satellites” problem

The satellites of the MW

Dark mattter subhalos in CDM



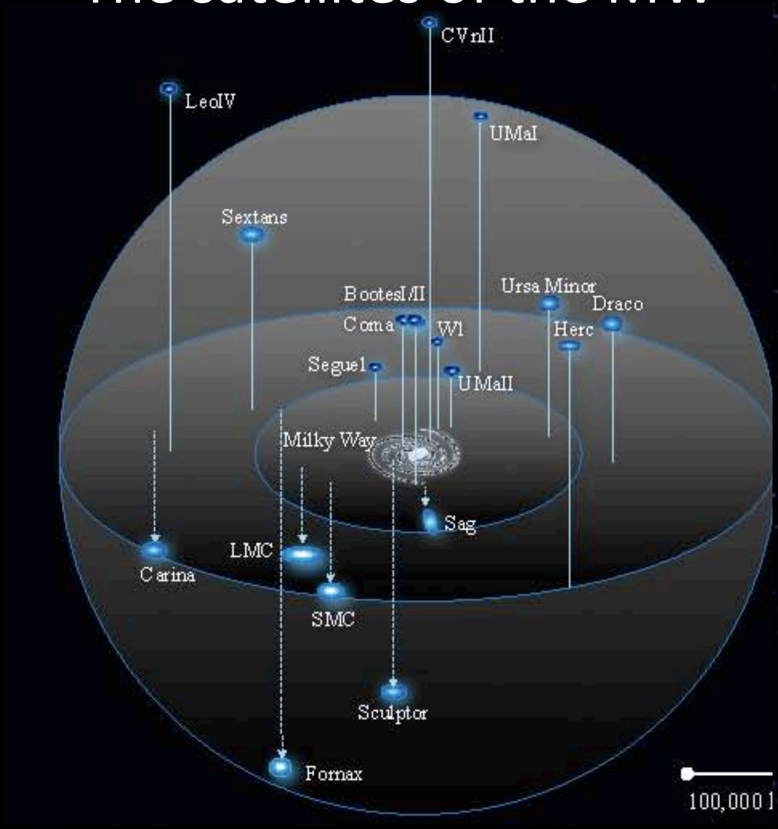
“Missing satellites” problem:

The Milky way has only about 25 satellites

BUT: CDM halos have a huge number of subhalos

The "missing satellites" problem

The satellites of the MW



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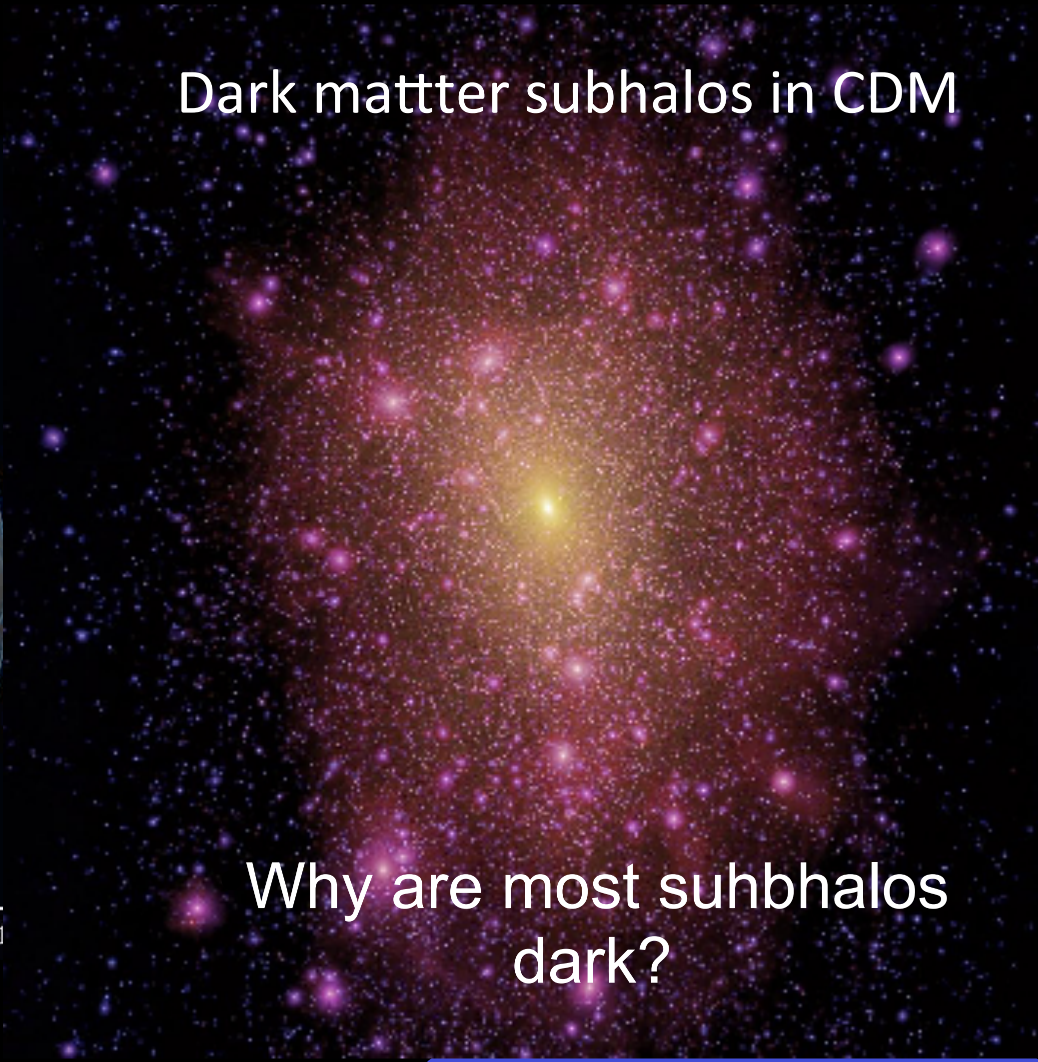
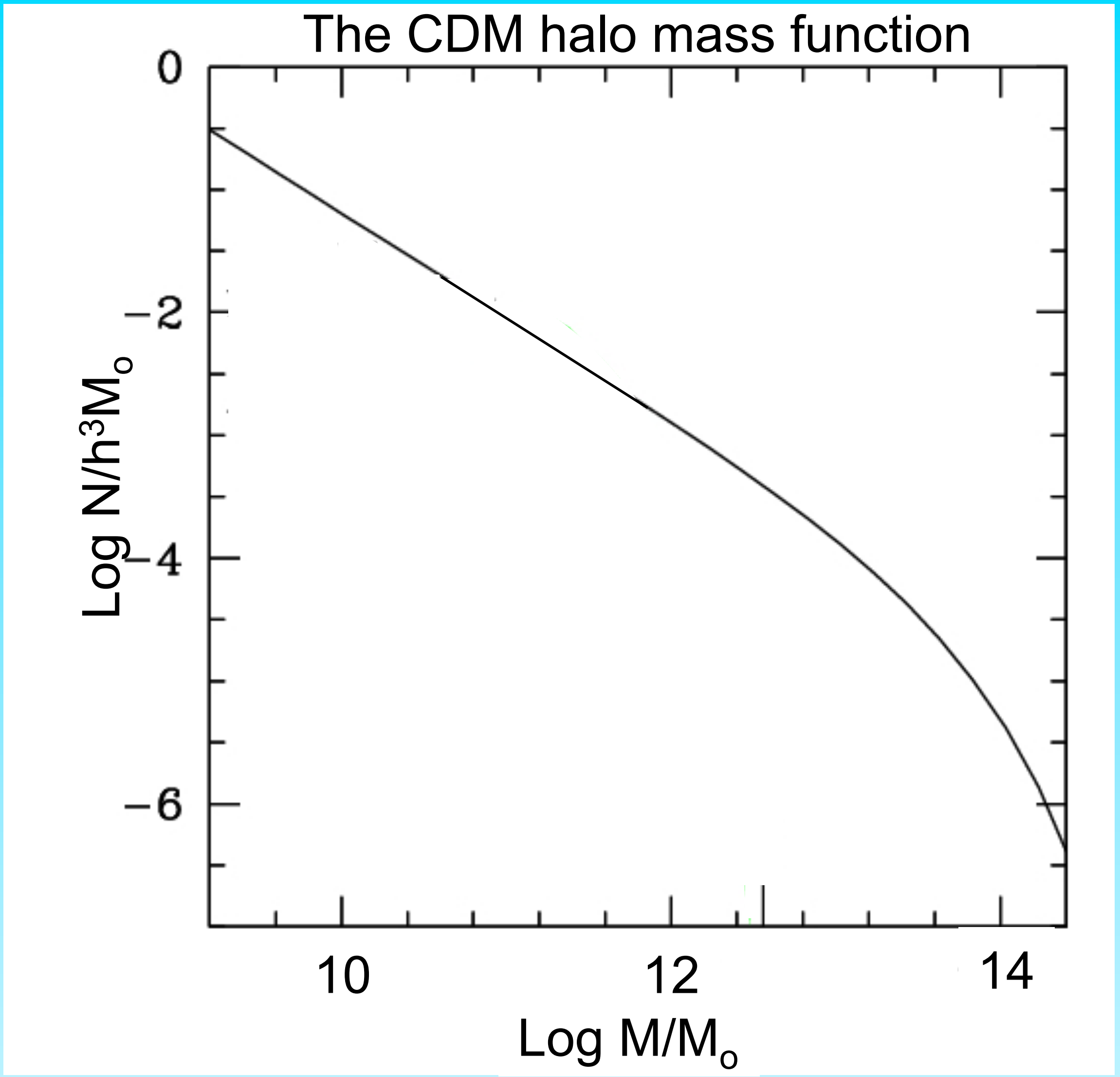


Why are most suhbhalos dark?

The CDM halo mass function

Jenkins et al. '01



The galaxy luminosity function

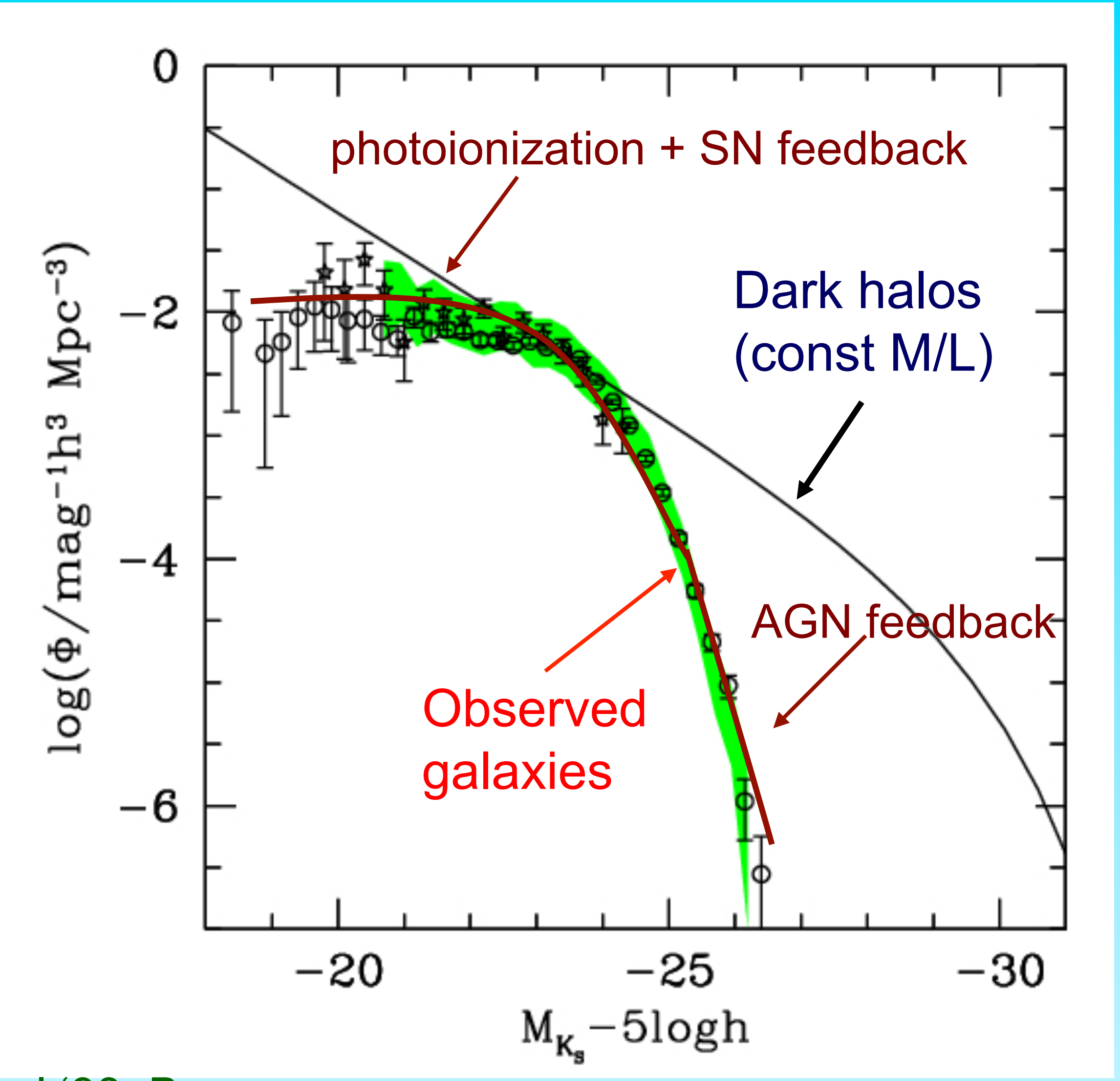
University of Durham

The halo mass function and the galaxy luminosity function have different shapes

⇨

Galaxy luminosity not just ∝ halo mass

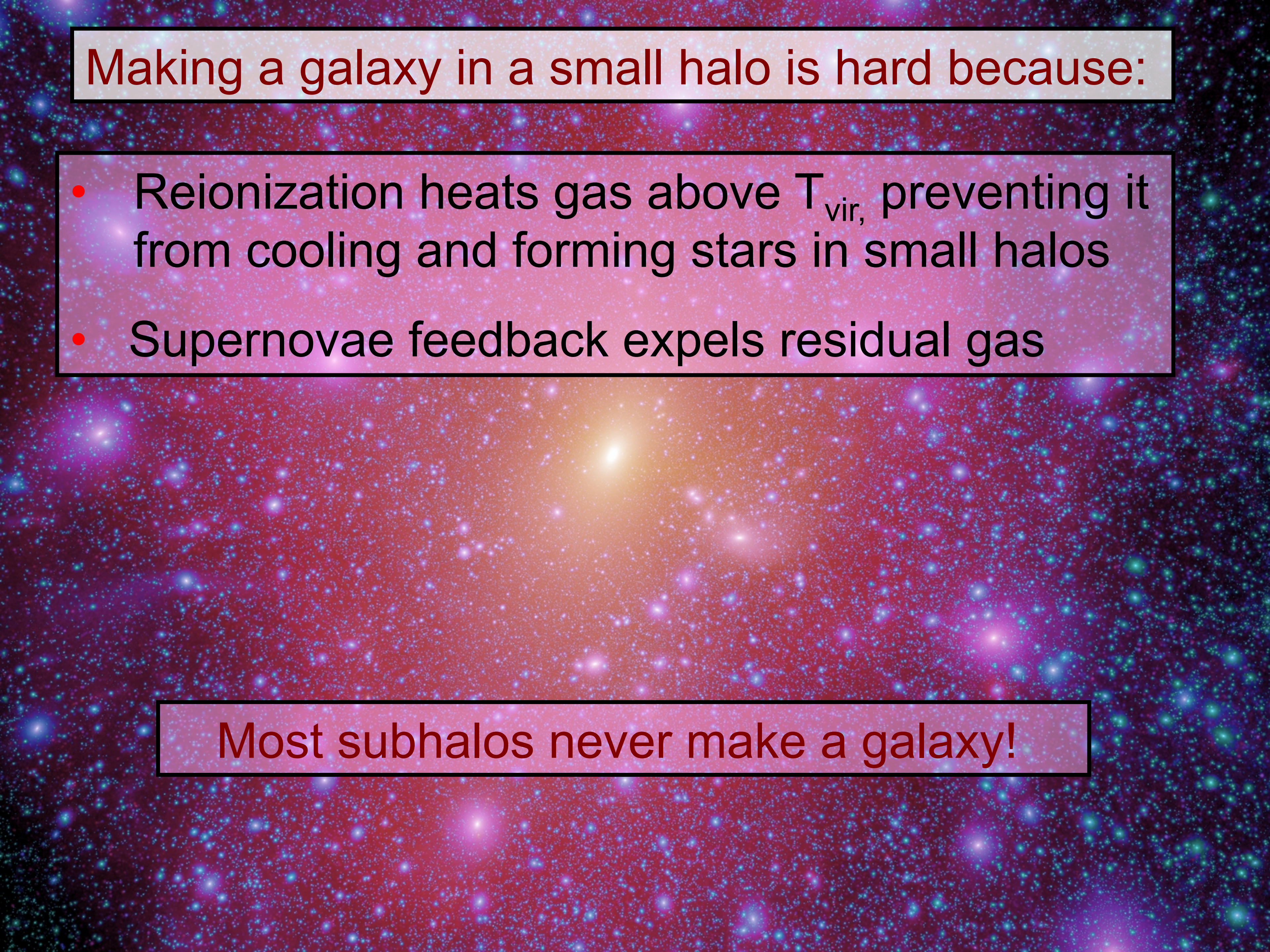
Complicated variation of M/L with halo mass



White & Frenk ‘91; Kauffmann et al ‘93; Benson et al ’03; Croton et al ‘06; Bower et al. ’06

Making a galaxy in a small halo is hard because:

- Reionization heats gas above $T_{vir,}$ preventing it from cooling and forming stars in small halos
- Supernovae feedback expels residual gas

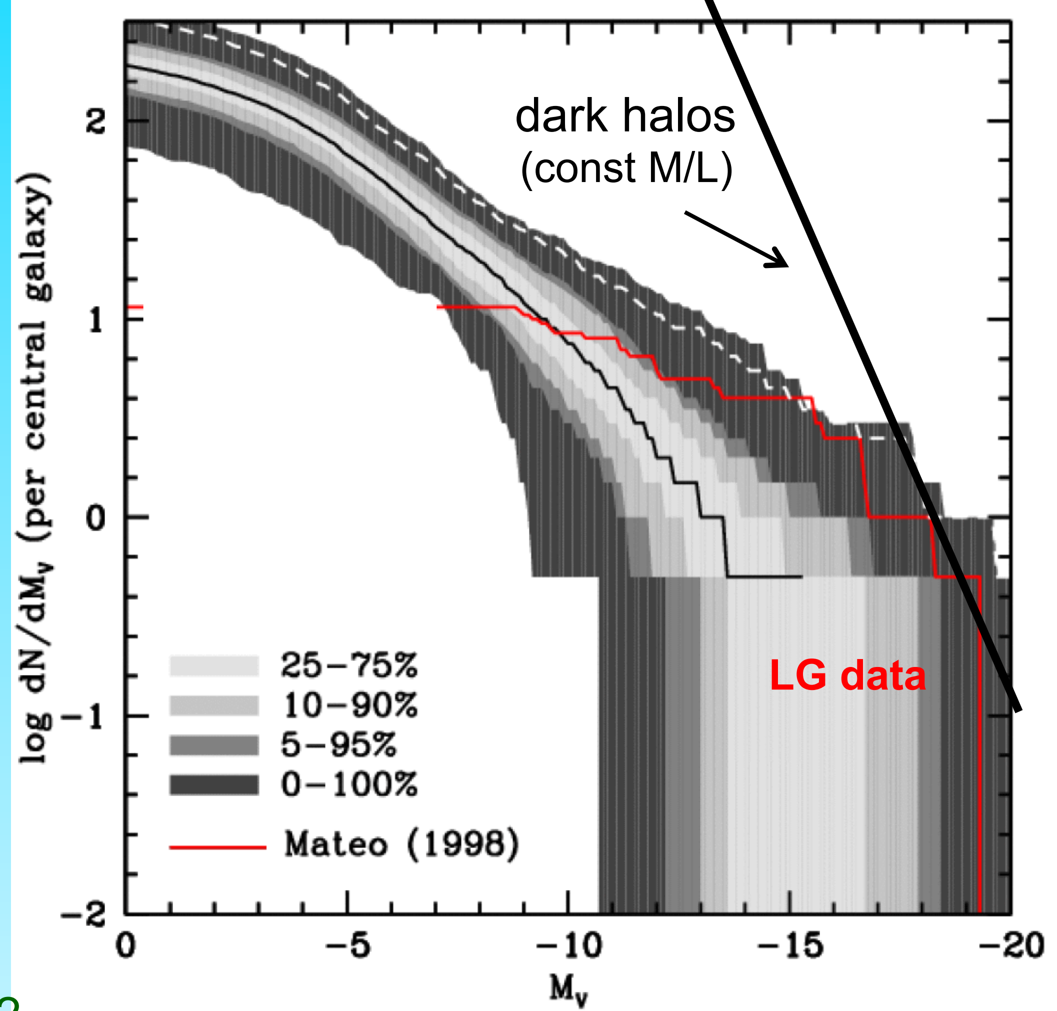


Most subhalos never make a galaxy!

Luminosity Function of Local Group Satellites

University of Durham

- Median model → correct abund. of sats brighter than M_V=-9 and V_{cir} > 12 km/s
- Model predicts many, as yet undiscovered, faint satellites
- LMC/SMC should be rare (~2% of cases)



Benson, Frenk, Lacey, Baugh & Cole ’02
(see also Kauffman etal ’93, Bullock etal ’01)

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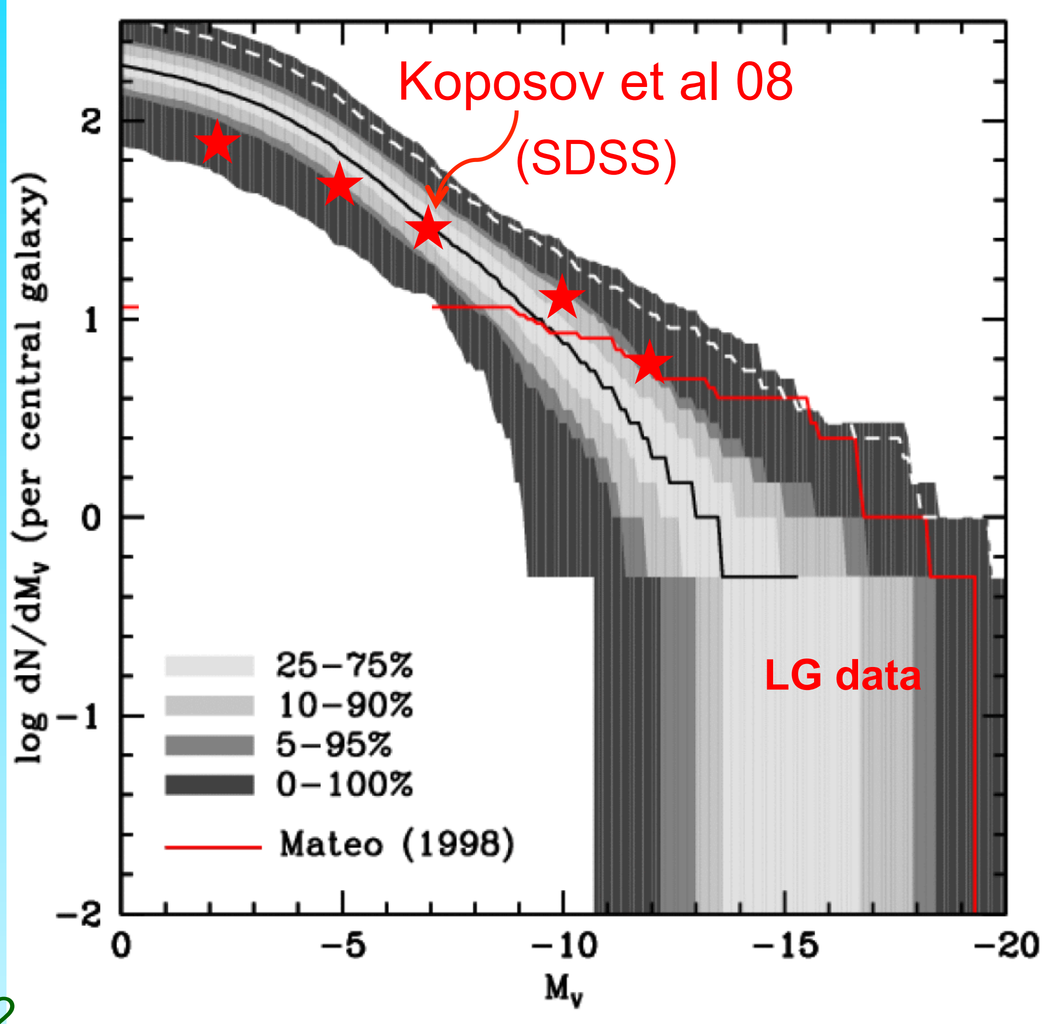


Benson, Frenk, Lacey, Baugh & Cole ’02
(see also Kauffman etal ’93, Bullock etal ’01)

EAGLE full hydro simulations

Local Group

Sawala et al ‘14

Far fewer satellite galaxies than CDM halos

EAGLE full hydro simulations

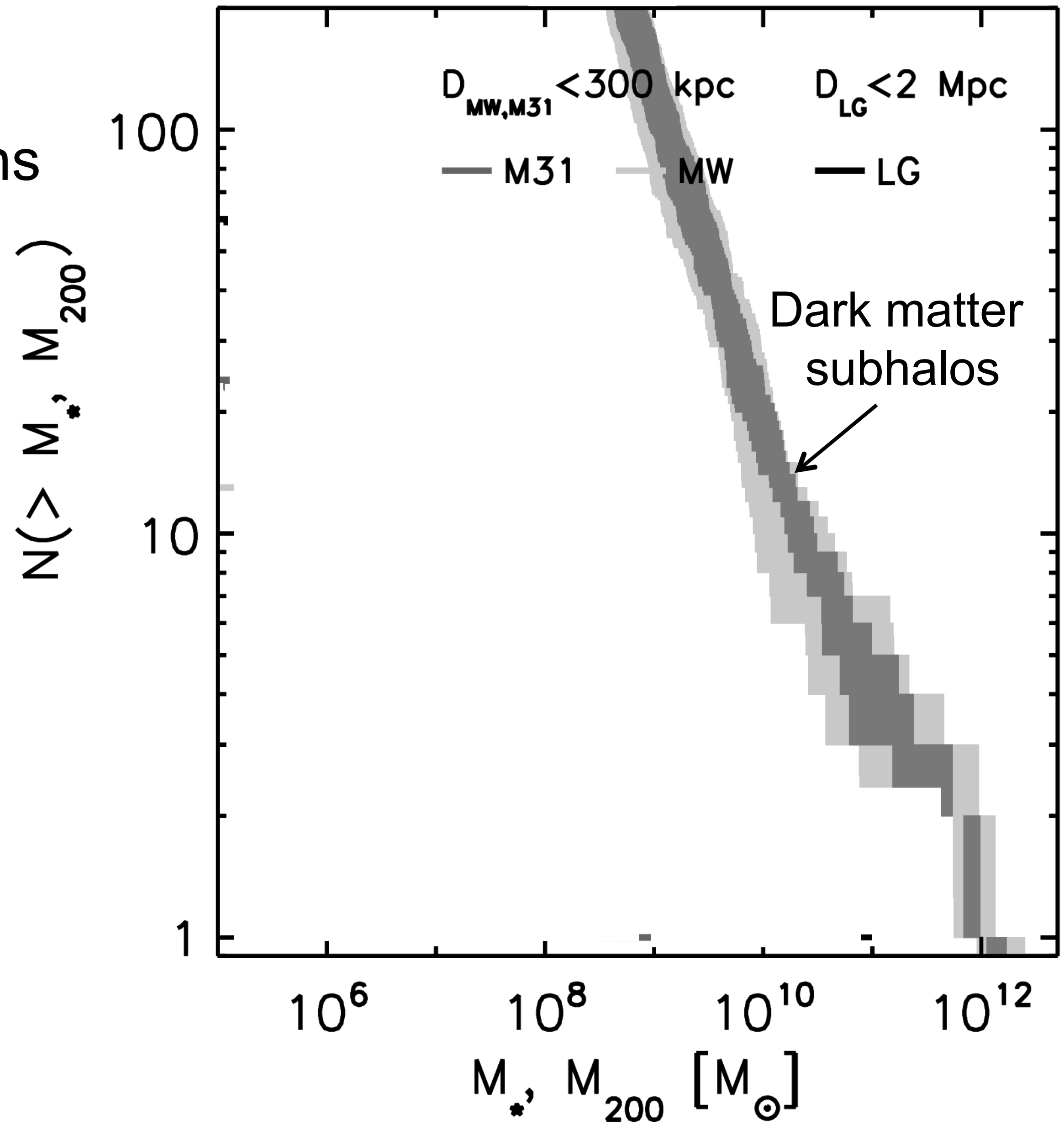
Local Group

Sawala et al '14

EAGLE Local Group simulation

University of Durham

Subhalo mass functions

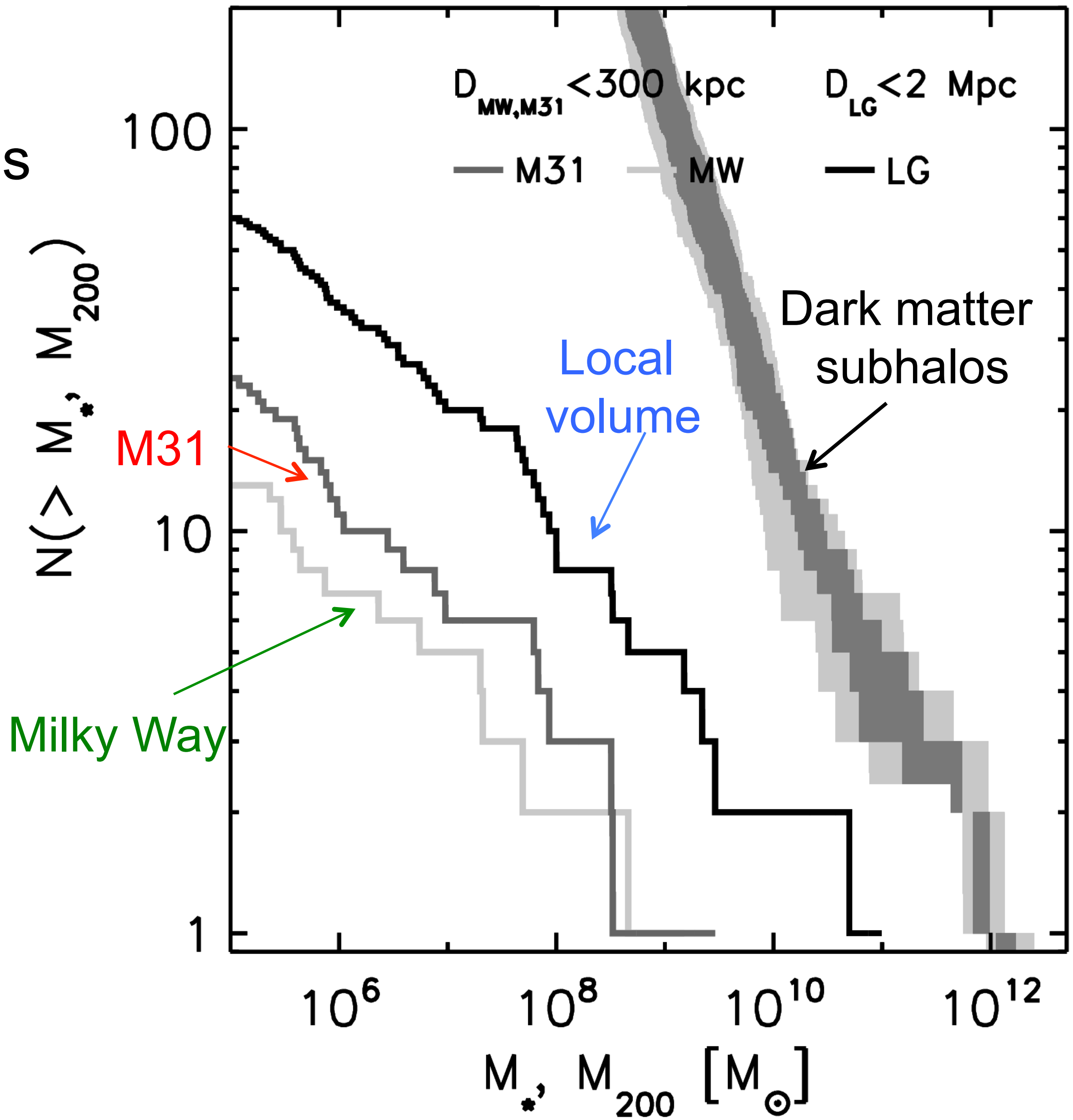


Sawala et al '14

EAGLE Local Group simulation

University of Durham

Stellar mass functions



Sawala et al '14

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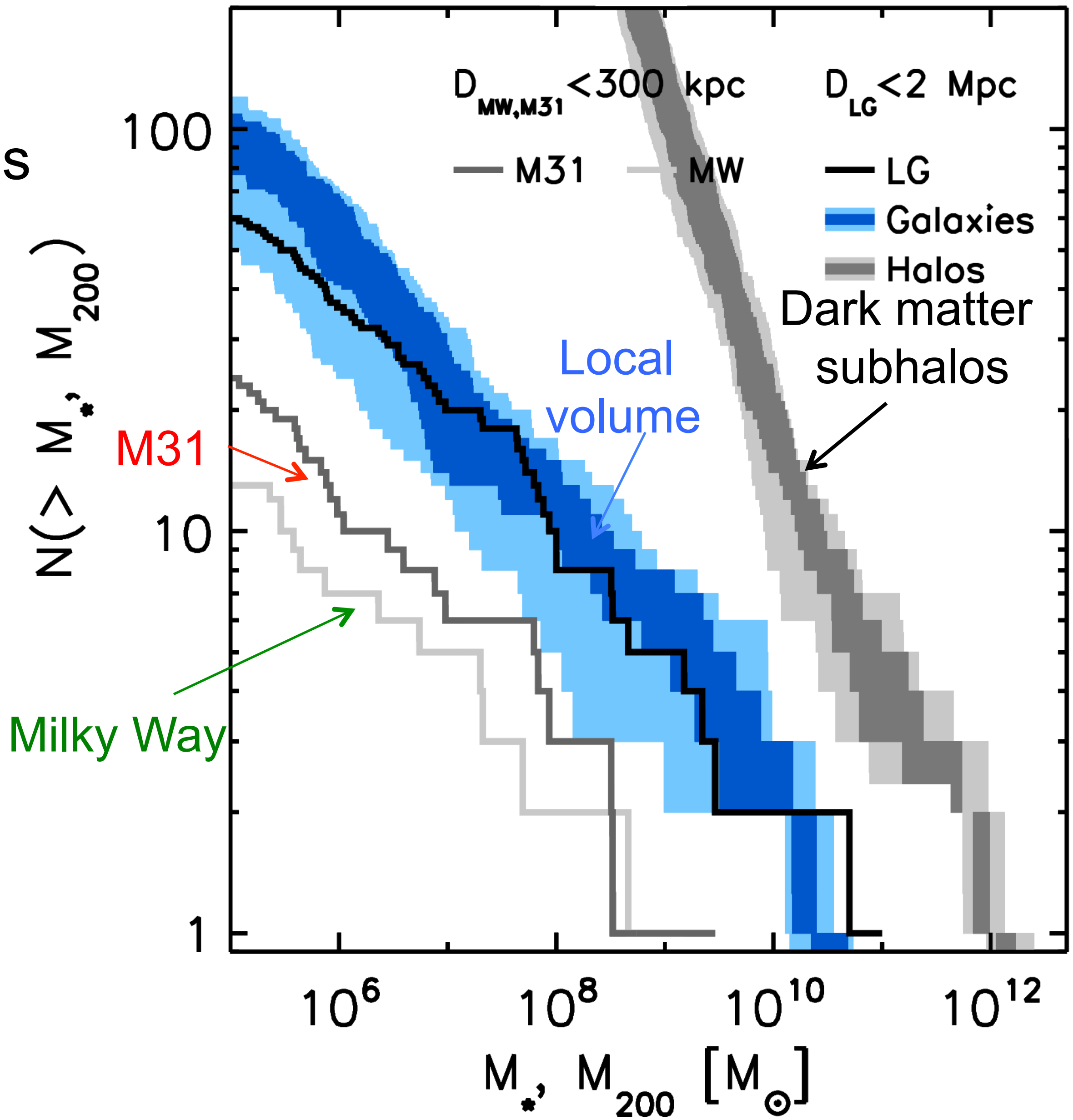
EAGLE Local Group simulation

Stellar mass functions



Sawala et al '14

EAGLE Local Group simulation

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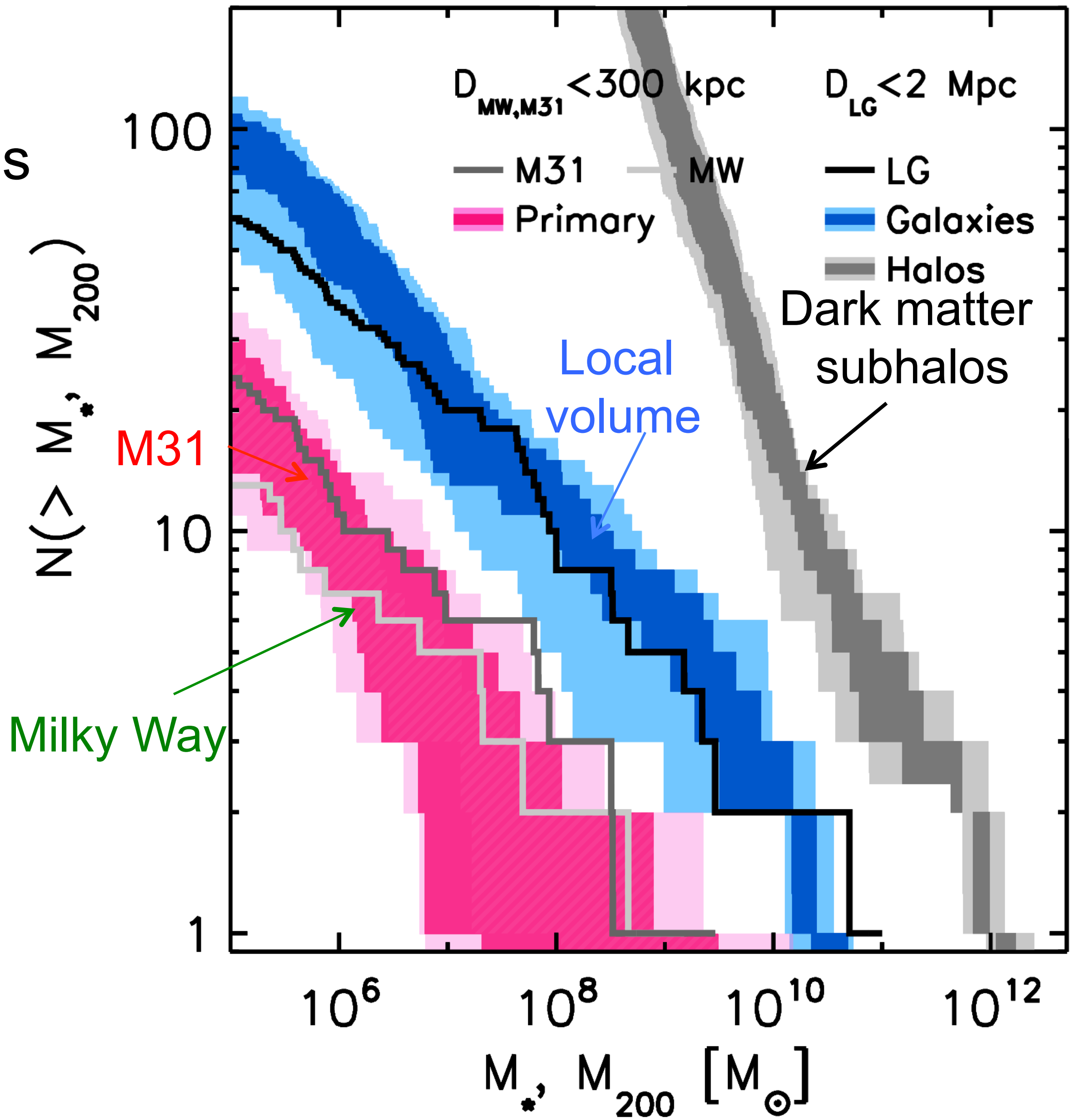
Stellar mass functions

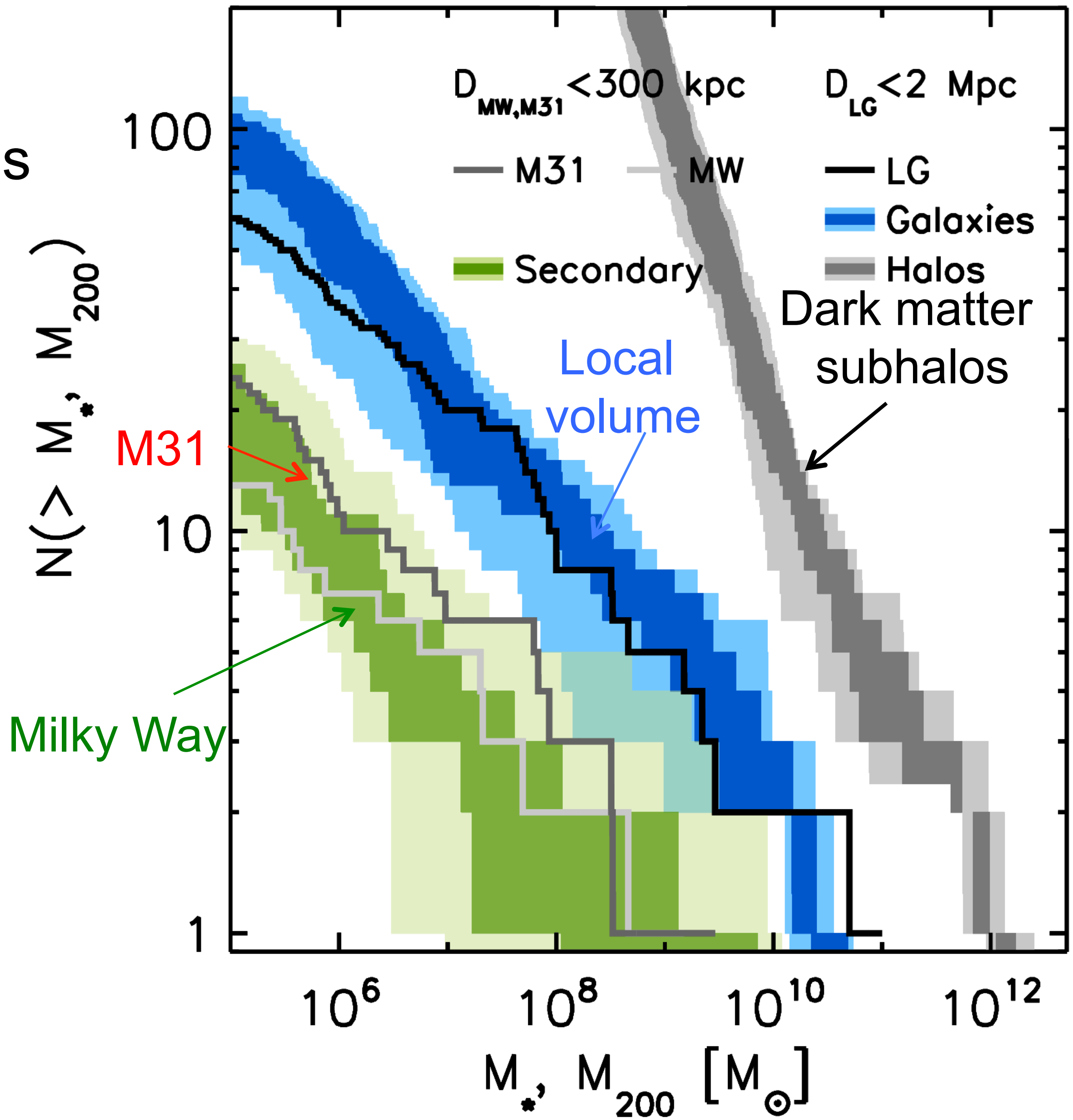


Sawala et al '14

EAGLE Local Group simulation

University of Durham

Stellar mass functions



Sawala et al '14

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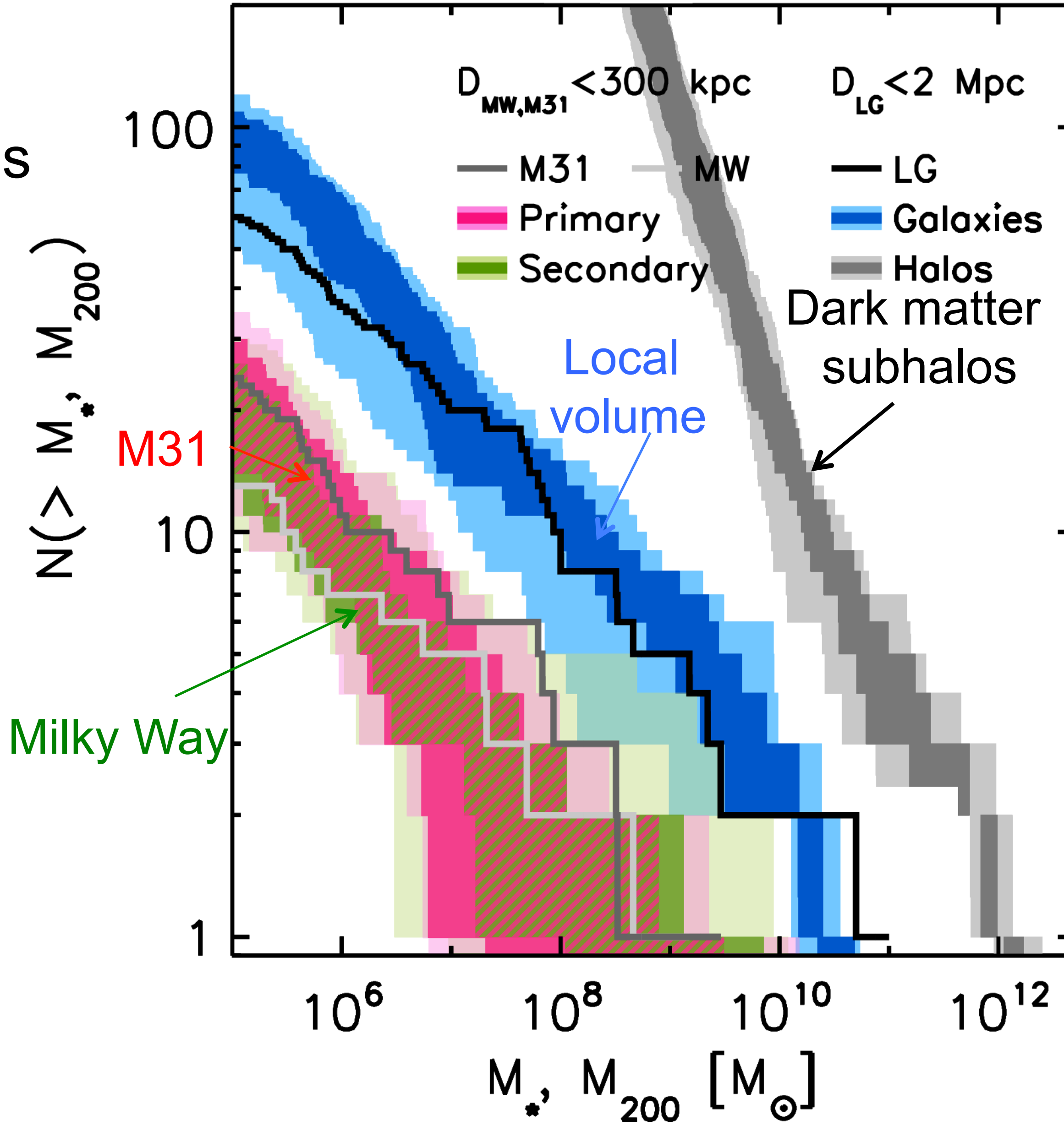
EAGLE Local Group simulation

Stellar mass functions



Sawala et al '14

ICC

Is there a “satellite problem” in CDM?

No, when galaxy formation is taken into account!

Four problems for CDM on small scales?

1. The “missing satellites” problem
2. The “too-big-to-fail” problem
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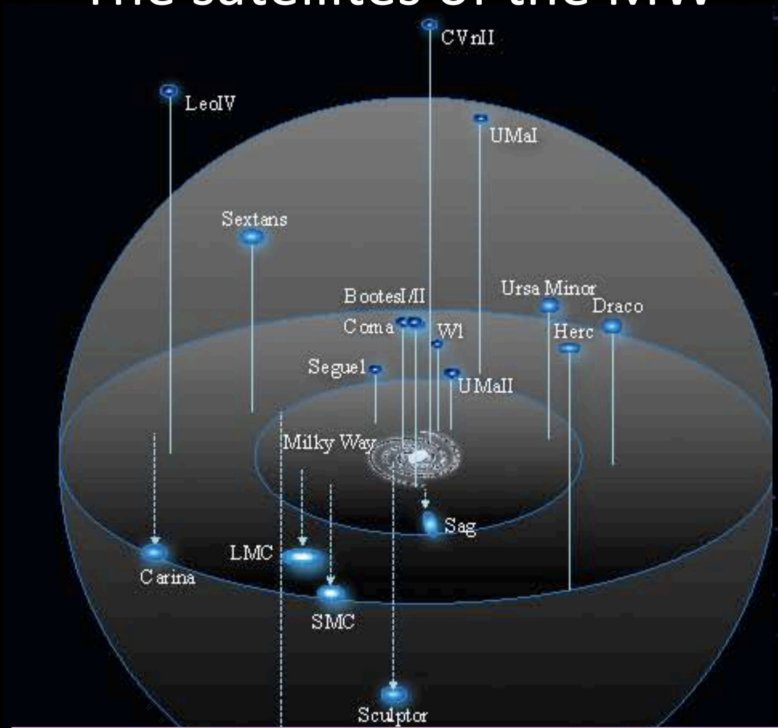
The “too-big-to-fail” problem

$$V_c = \sqrt{\frac{GM}{r}}$$

$$V_{max} = \max V_c$$

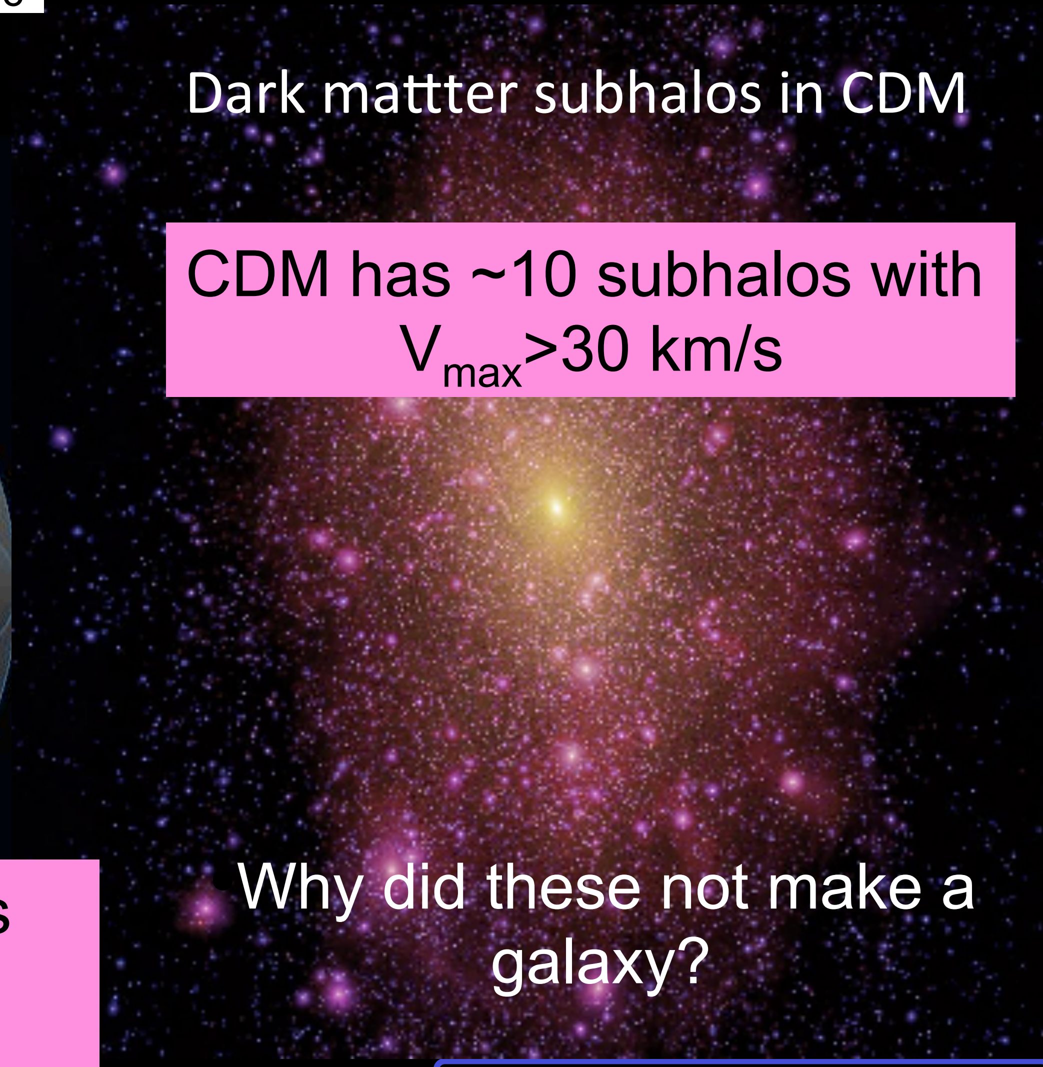
The satellites of the MW



MW has only 3 satellites with V_{max}>30 km/s

(LMC, SMC, Sgr)

Dark mattter subhalos in CDM



CDM has ~10 subhalos with V_{max}>30 km/s

Why did these not make a galaxy?

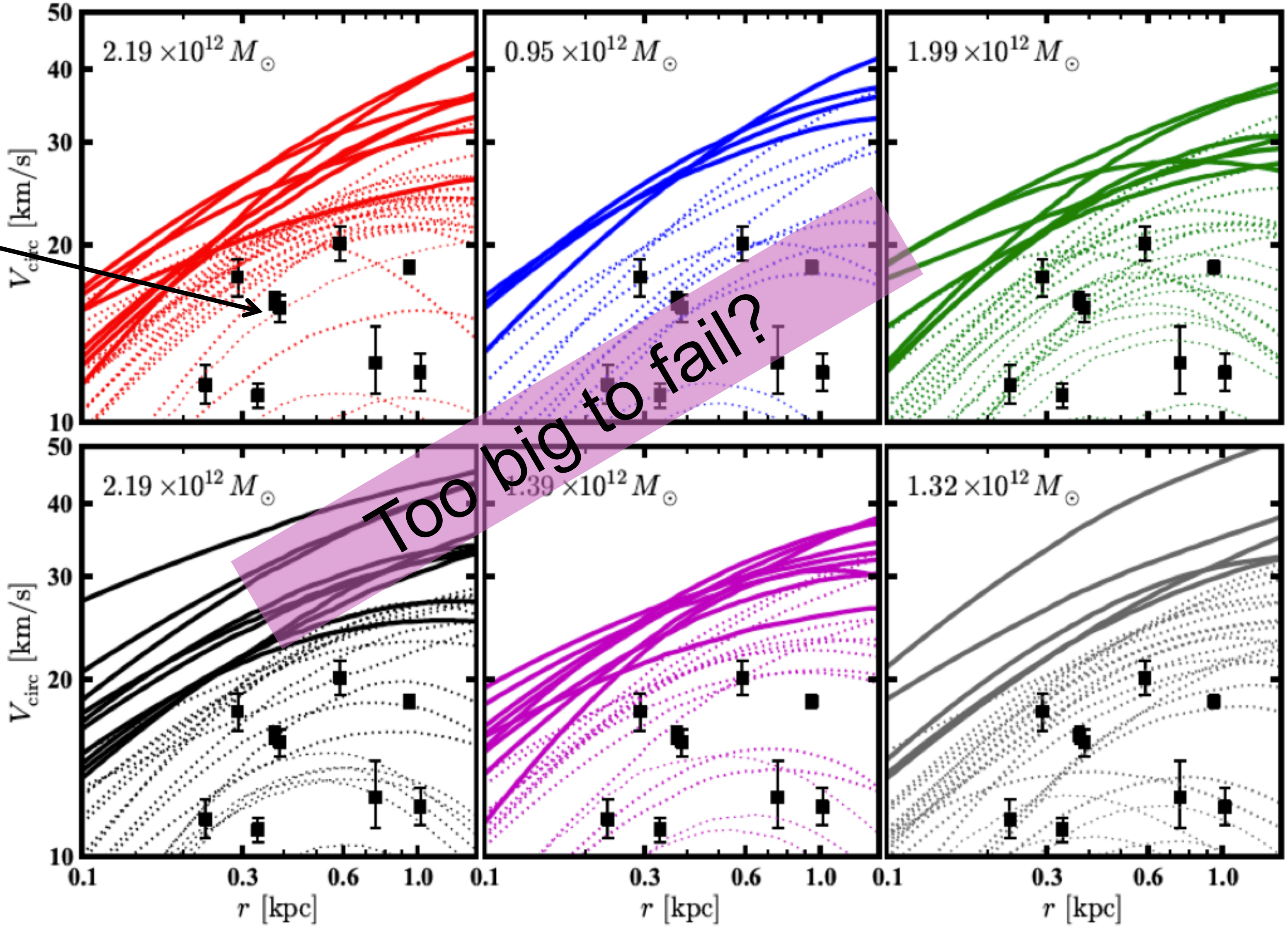
Rotation curves of Aquarius subhalos

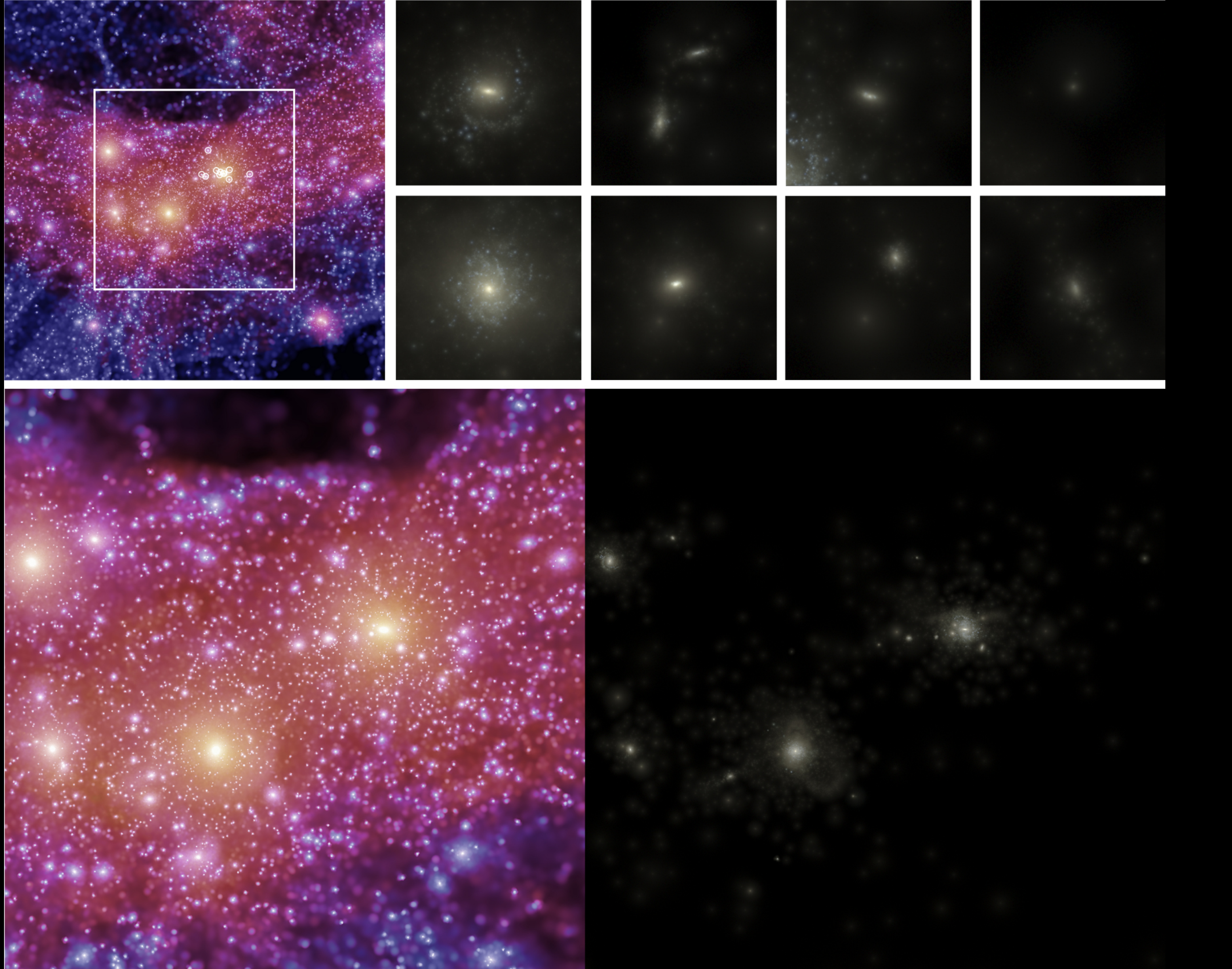
Boylan-Kolchin et al. '11

$$V_c = \sqrt{\frac{GM}{r}}$$

9 dwarf satellites of Milky Way: mass within half-light radius

Excludes LMC, SMC, Sagittarius





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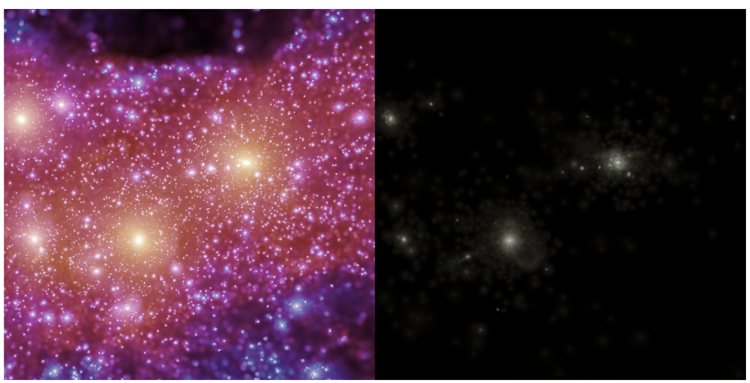
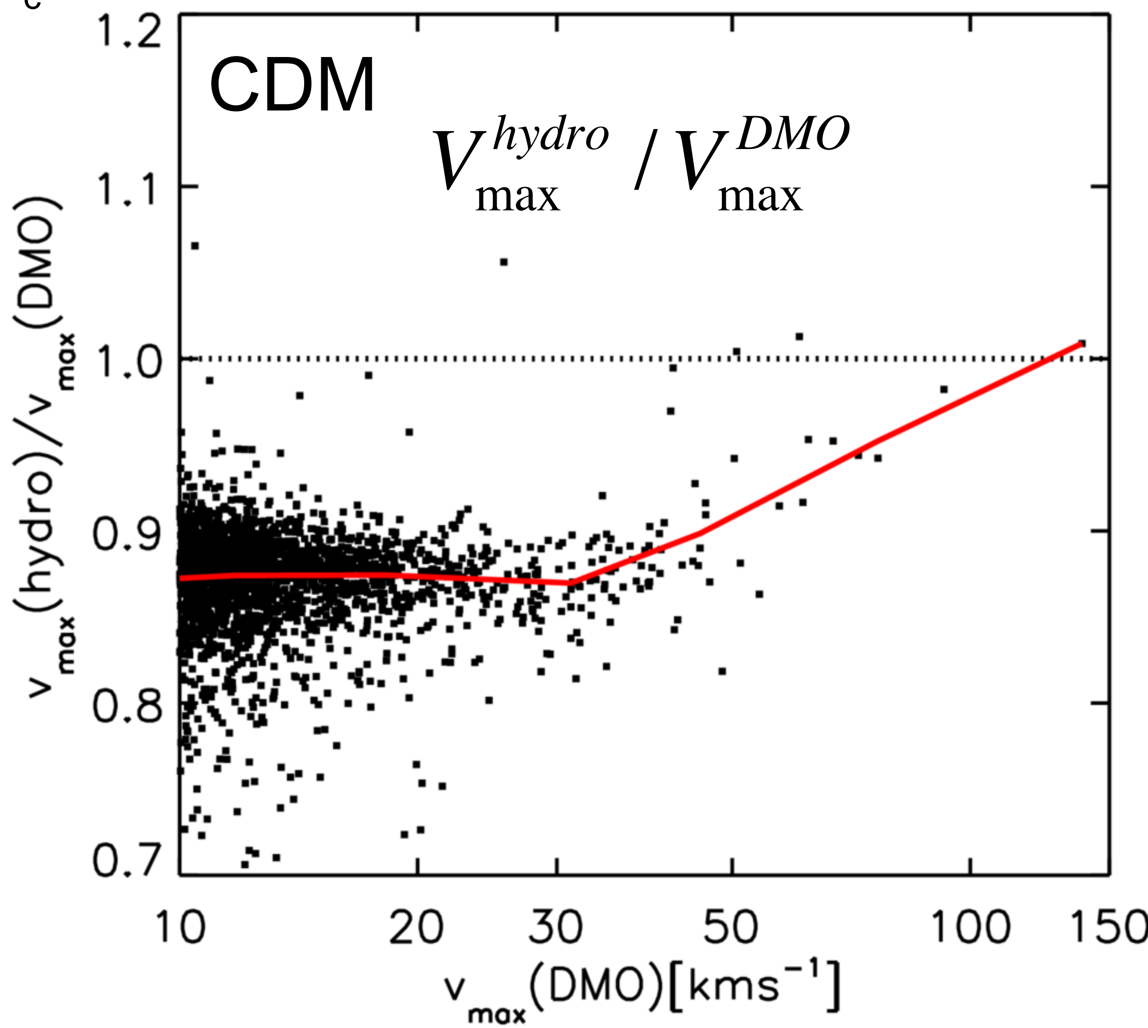
To-big-to-fail in CDM: baryon effects

$$V_c = \sqrt{\frac{GM}{r}} \qquad V_{max} = \max V_c$$

Reduction in V_{max} due to SN feedback:

→ Lowers halo mass & thus halo growth rate





Sawala et al. ‘13, ‘14

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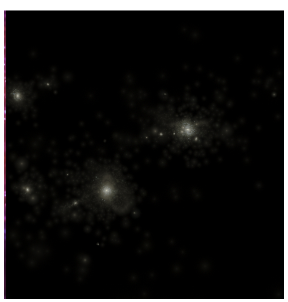
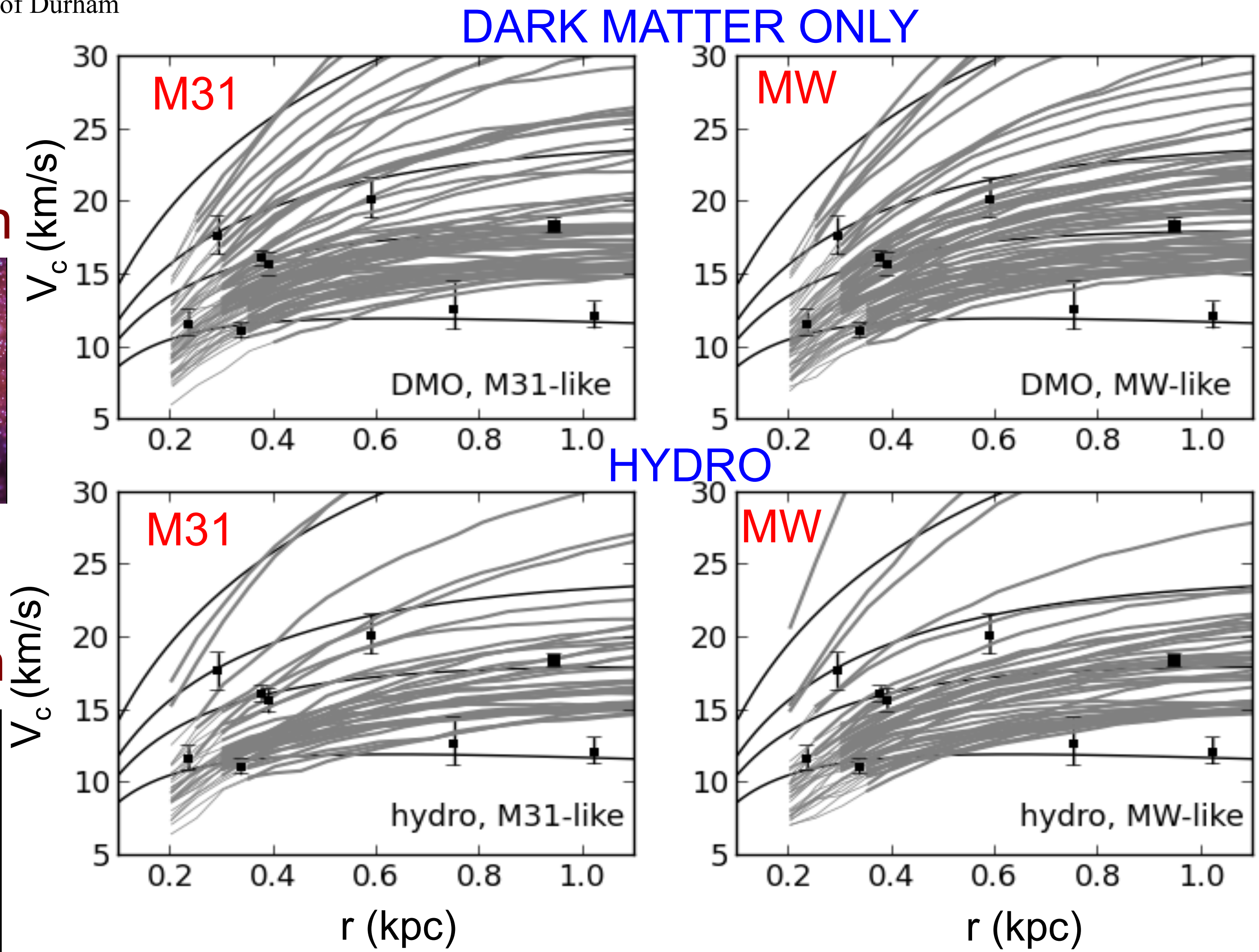
Too-big-to-fail: the baryon bailout

DM-only simulation

Gas simulation





Sawala et al. '14

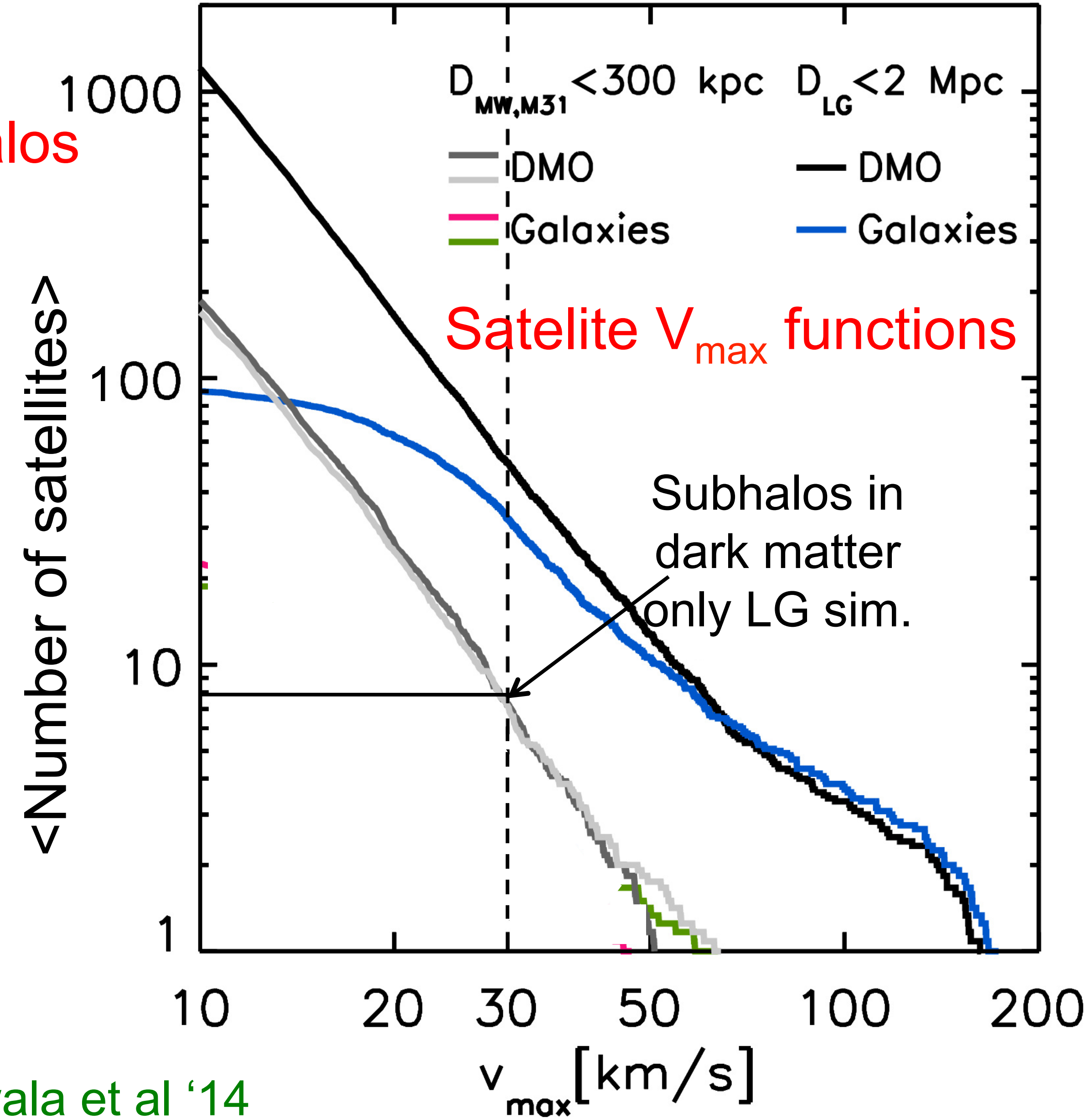
Number of subhalos of given V_{max} is greatly reduced in gas simulations

Institute for Computational Cosmology

Too-big-to-fail: the baryon bailout

University of Durham

DM only sims → ~10 halos with V_{max}> 30 km/s

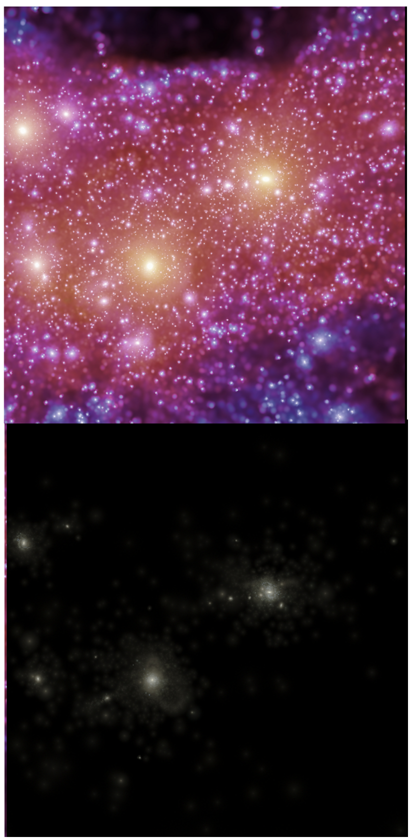


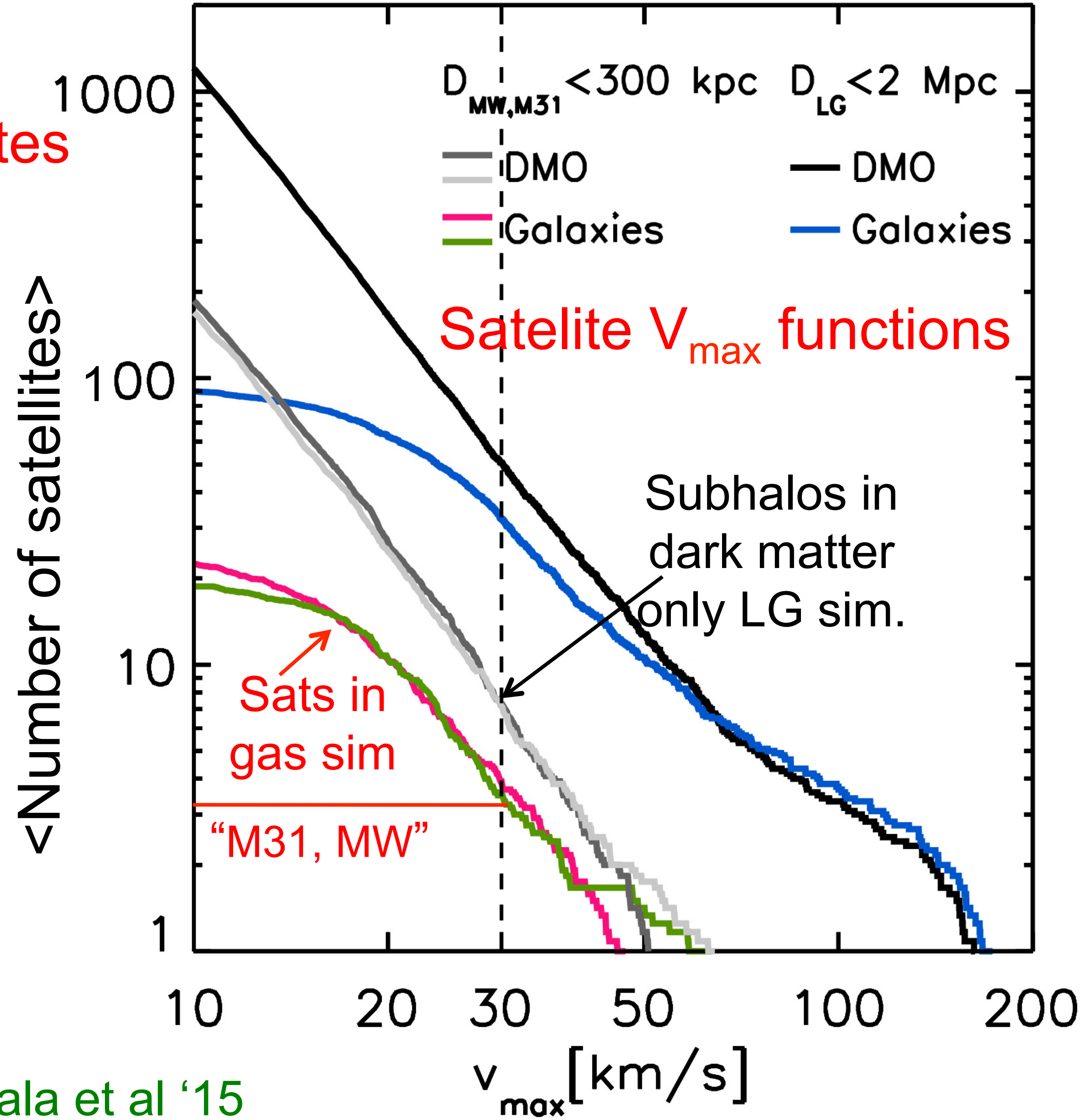
Sawala et al '14

University of Durham

Too-big-to-fail: the baryon bailout

Hydro sims → ~3 satellites with V_{max}> 30 km/s



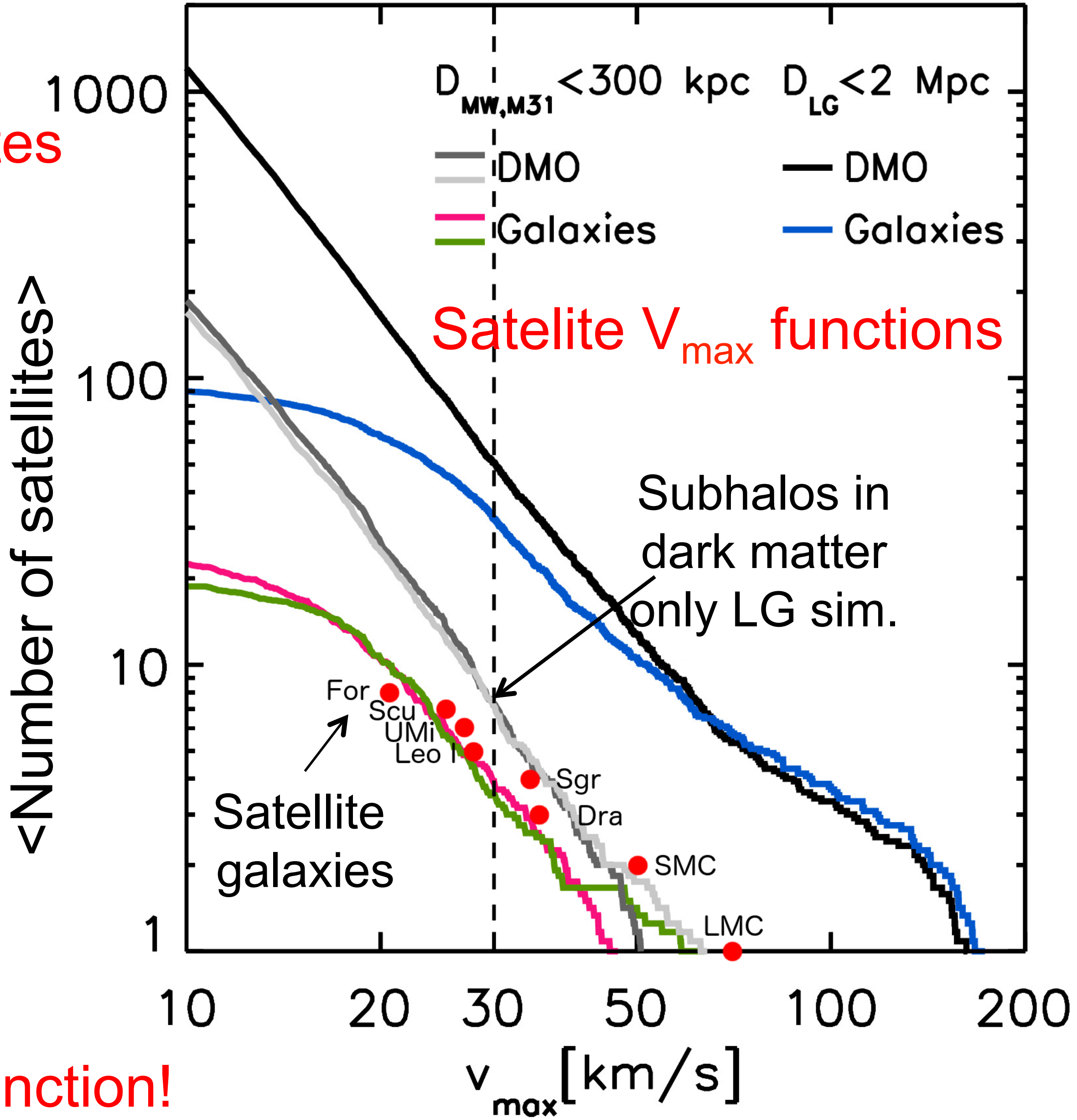


Sawala et al '15

Too-big-to-fail: the baryon bailout

University of Durham

Hydro sims → ~3 satellites with V_{max}> 30 km/s



. and with correct V_{max} function!

Is there a “too-big-to-fail” problem in CDM?

No, when galaxy formation is taken into account!

Four problems for CDM on small scales?

1. The "missing satellites" problem
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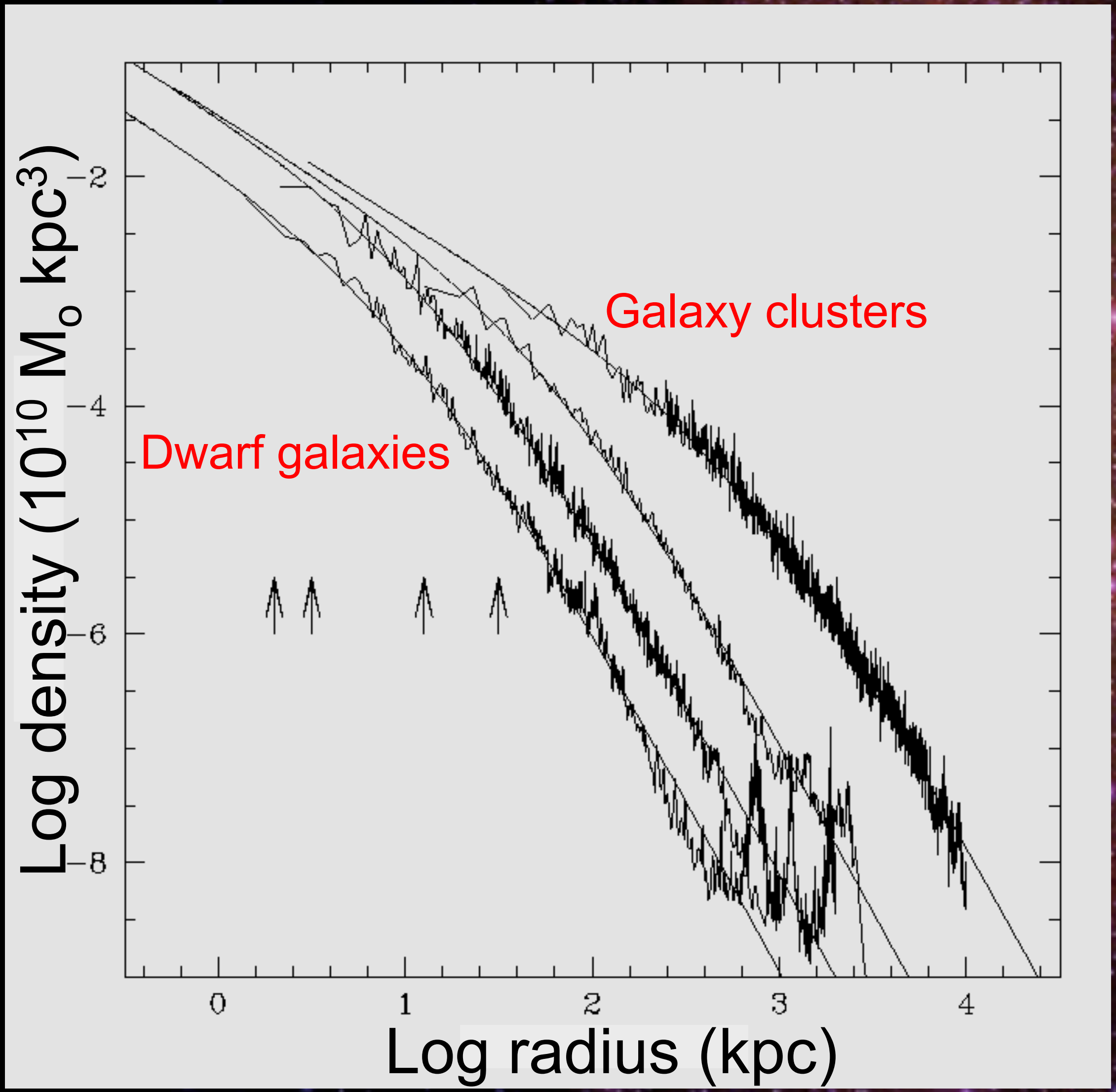
A challenge to Andi Burkert

The Density Profile of Cold Dark Matter Halos



Shape of halo profiles ~independent of halo mass & cosmological parameters

Density profiles are “cuspy” - no `core’ near the centre

Fitted by simple formula:

$$\frac{\rho(r)}{\rho_{crit}} = \frac{\delta_c}{(r/r_s)(1+r/r_s)^2}$$

(Navarro, Frenk & White ’97)

More massive halos and halos that form earlier have higher densities (bigger δ)

The core-cusp problem

"Core-cusp" problem:

CDM & subhalos have cuspy profiles

BUT: kinematical data are said to "show" that the dwarf satellites of the Milky Way have cores


Dwarf galaxies in Eagle have NFW cusps!
Sawala et al '15

Dwarf galaxies around the Milky Way

Many claims that dwarf spheroidal satellites have density cores

- e.g. Gilmore et al. '07, Kuzio de Naray '08 and many more

Fornax

Sculptor

Leo I

Sextans

Carina

Sagittarius

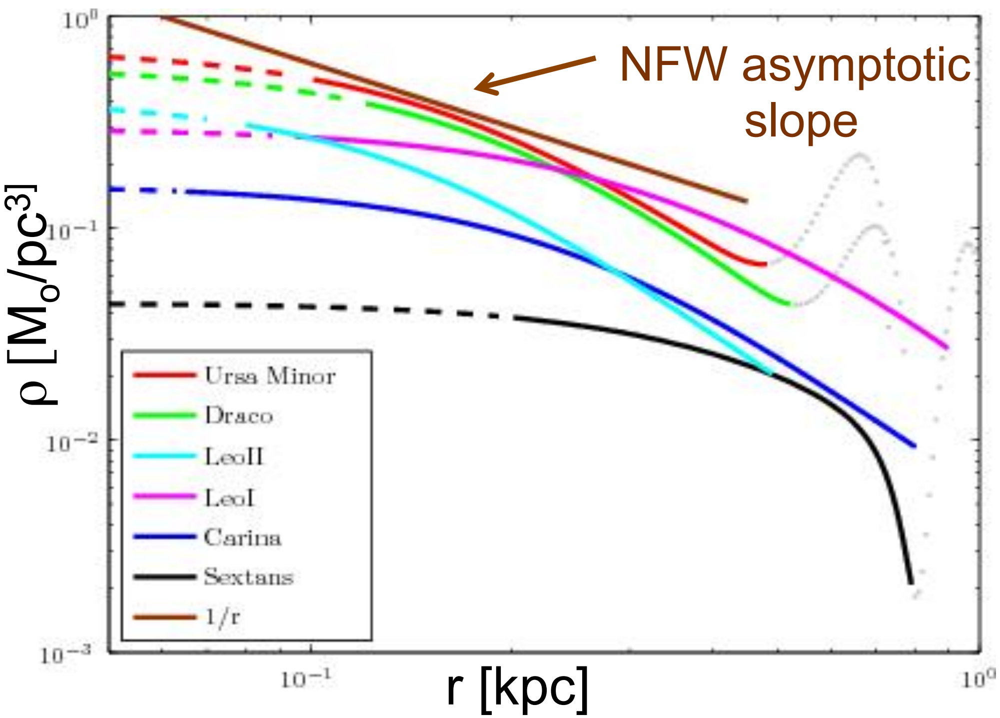
Evidence for warm dark matter?

University of Durham

THE ASTROPHYSICAL JOURNAL, 663:948–959, 2007 July 10

THE OBSERVED PROPERTIES OF DARK MATTER ON SMALL SPATIAL SCALES

GERARD GILMORE,[1] MARK I. WILKINSON,[1,2] ROSEMARY F. G. WYSE,[3] JAN T. KLEYNA,[4] ANDREAS KOCH,[5,6]
N. WYN EVANS,[1] AND EVA K. GREBEL[6,7]



Inferred density profiles for 6 dwarf spheroidals

"…dark matter forms cored mass distributions, with a core scale length of greater than about 100pc, and always has a maximum central density in a narrow range…"

"…(keV) sterile neutrino particles have been discussed as relevant in just the spatial and density range we have derived here."

ICC University of Durham

Fits assuming NFW →

Dwarf sphs: cores or cusps?

Jeans eqn:

$$\frac{GM(r)}{r} = -\sigma_r^2 \left[\frac{d\ln\rho_*}{d\ln r} + \frac{d\ln\sigma_r^2}{d\ln r} + 2\beta \right]$$

from Aquarius sim

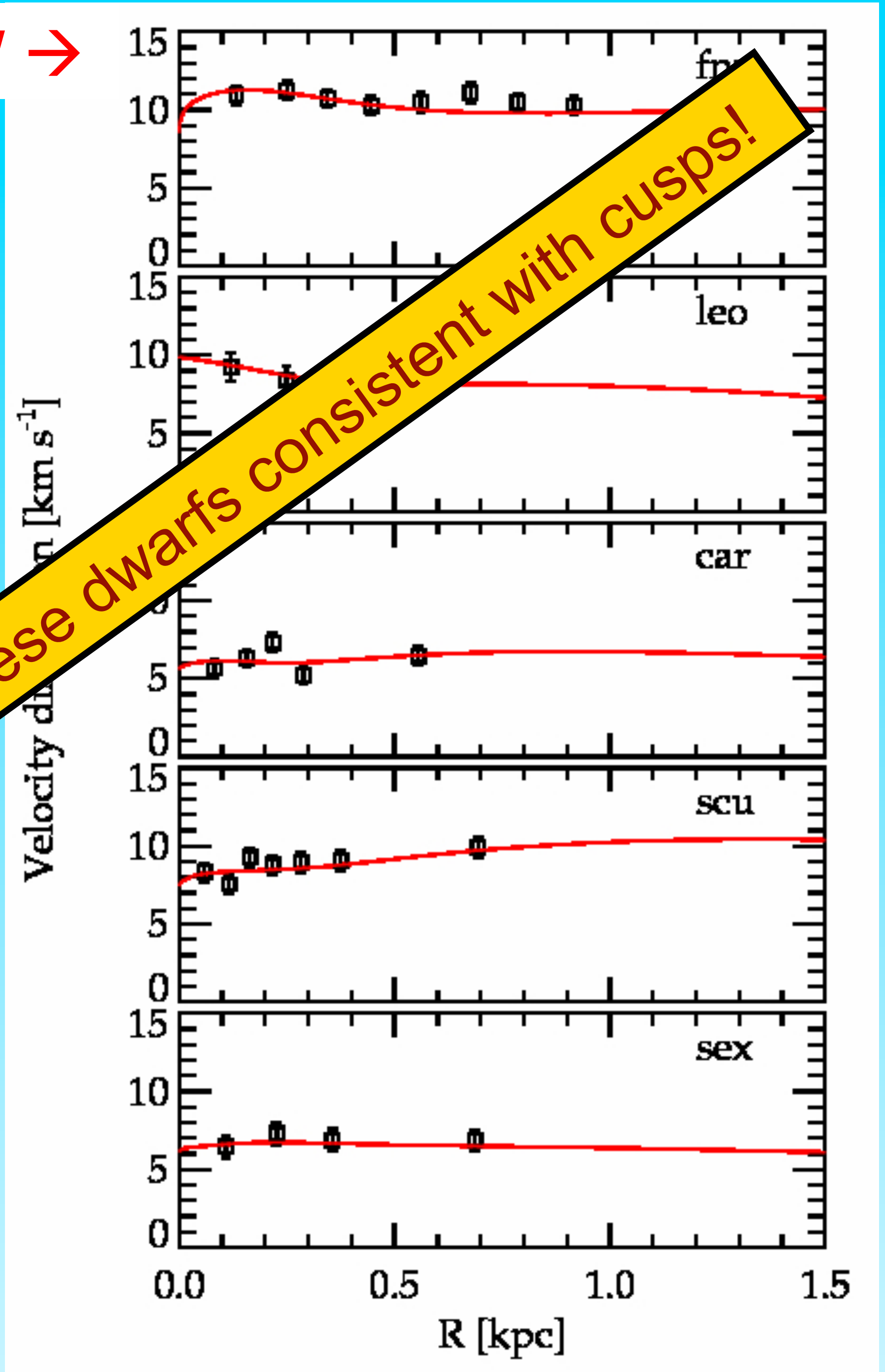
Cuspy!

vel. an... py

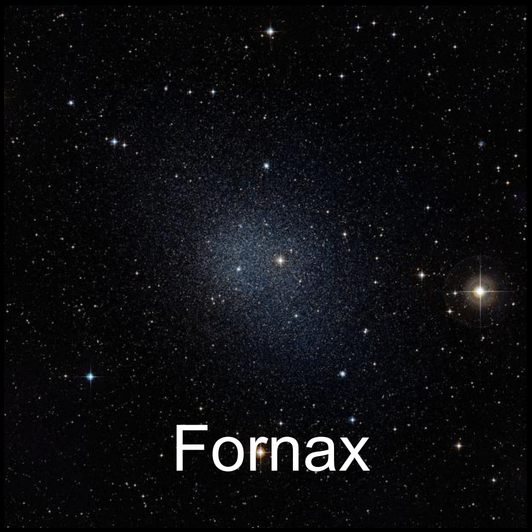
- Assume isotropi... ts
- Solve for σ...
- Com... with observed σ_r (r)
- ... best fit" subhalo

Photometric and kinematical data for these dwarfs consistent with cusps!

Strigari, Frenk & White '10



Dwarf galaxies around the Milky Way



Fornax

Sculptor

Leo I

Sextans

Carina

Sagittarius

The DM halo of the Sculptor dwarf

University of Durham

Sculptor has two stellar pops:

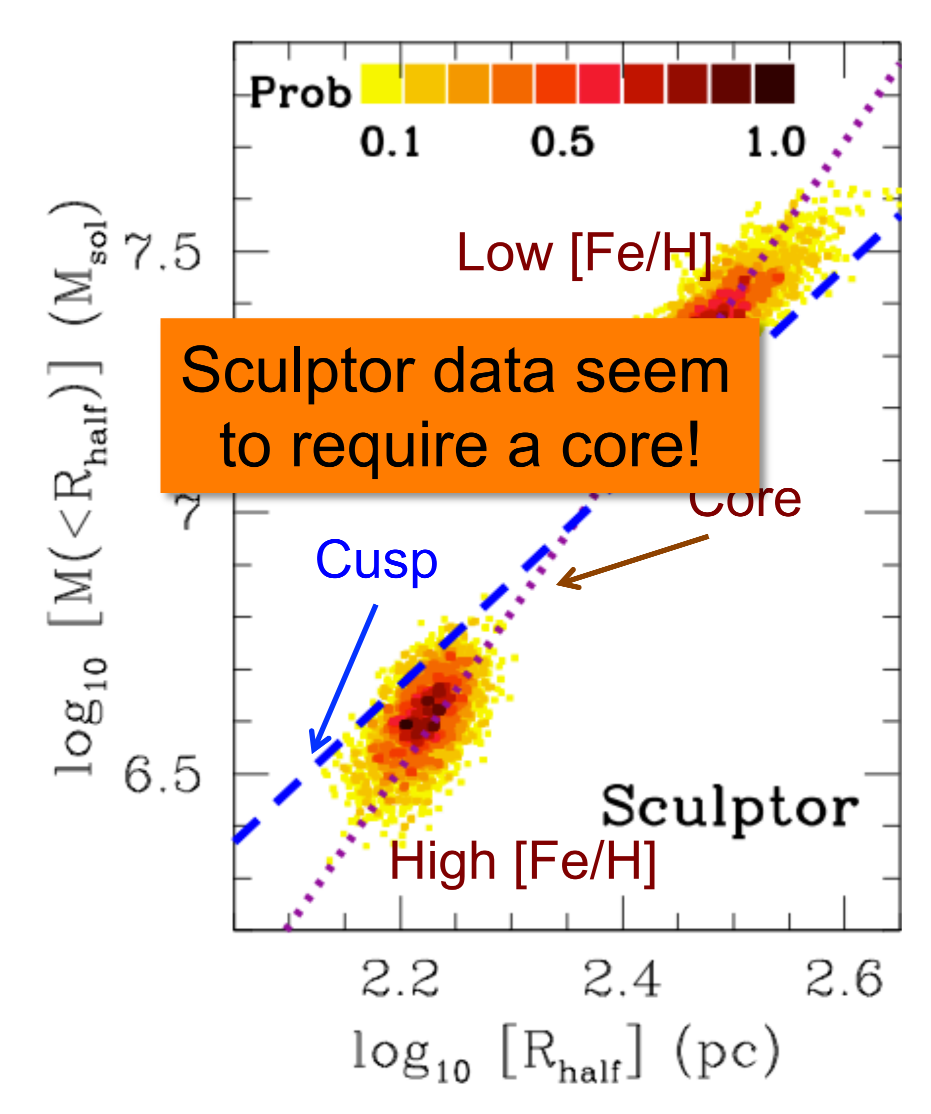
(i) centrally concentrated, high [Fe/H]

(ii) extended, low [Fe/H]

$$M(<r) = \mu \frac{r < \sigma_{los}^{2} >}{G}$$

Walker '10; Wolf et al '10→

if $r=r_{1/2}$, $\mu=2.5$, independently of model assumptions!



Walker & Peñarrubia '11

University of Durham

The DM halo of the Sculptor dwarf

Strigari, Frenk & White '14

Distribution function analysis of 2 metallicity pop. data of Battaglia et al.

Assume pops in equil. in NFW halo: $$\rho(r) = \frac{\rho_s}{x(1+x)^2}$$

For each population: $$f(E,J) = g(J)h(E),$$

Parametrize: $$g(J) = \left[\left(\frac{J}{J_\beta}\right)^{\frac{b_0}{\alpha}} + \left(\frac{J}{J_\beta}\right)^{\frac{b_1}{\alpha}}\right]^{\alpha}$$

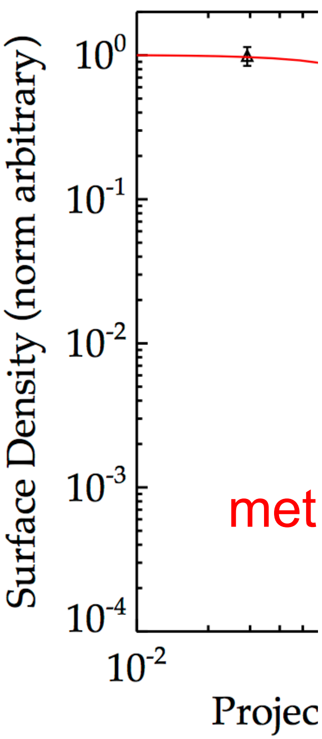
$$h(E) = \begin{cases} NE^a(E^q + E_c^q)^{d/q}(\Phi_{lim} - E)^e & \text{for } E < \Phi_{lim} \\ 0 & \text{for} E \geq \Phi_{lim}, \end{cases}$$

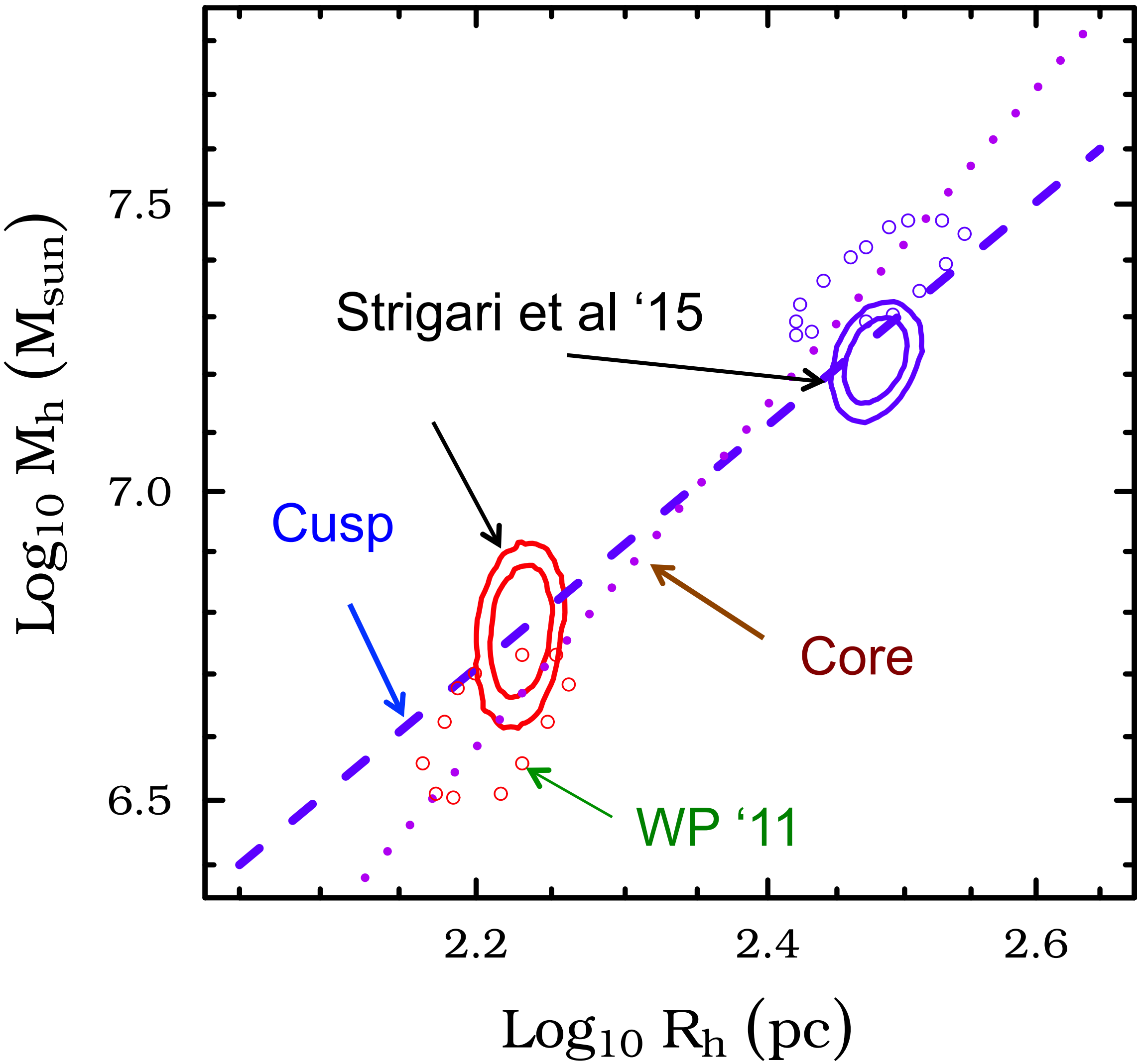
Find best-fit parameters using MCMC

The DM halo of the Sculptor dwarf

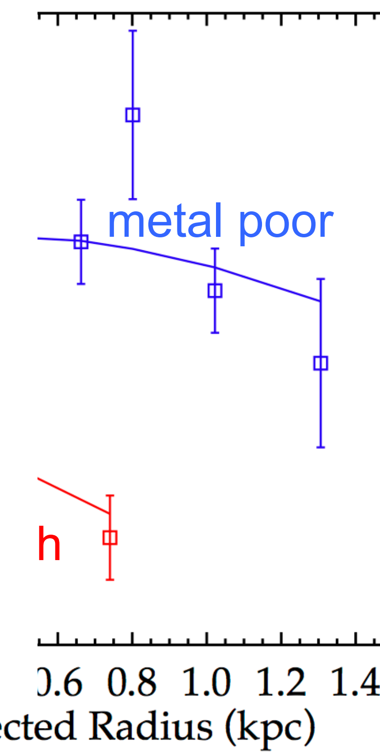
University of Durham

Distribution function analysis





y dispersion



Data cons

eq

Strigari, Fre...

Cores or cusps in the dwarf sph. satellites of the MW?

When sufficiently general models are considered, even best kinematical data cannot distinguish cores from NFW cusps in the dwarf spheroidal satellites of the Milky Way

How about in field dwarf galaxies?

Data are not as detailed, but some dwarfs have disks:

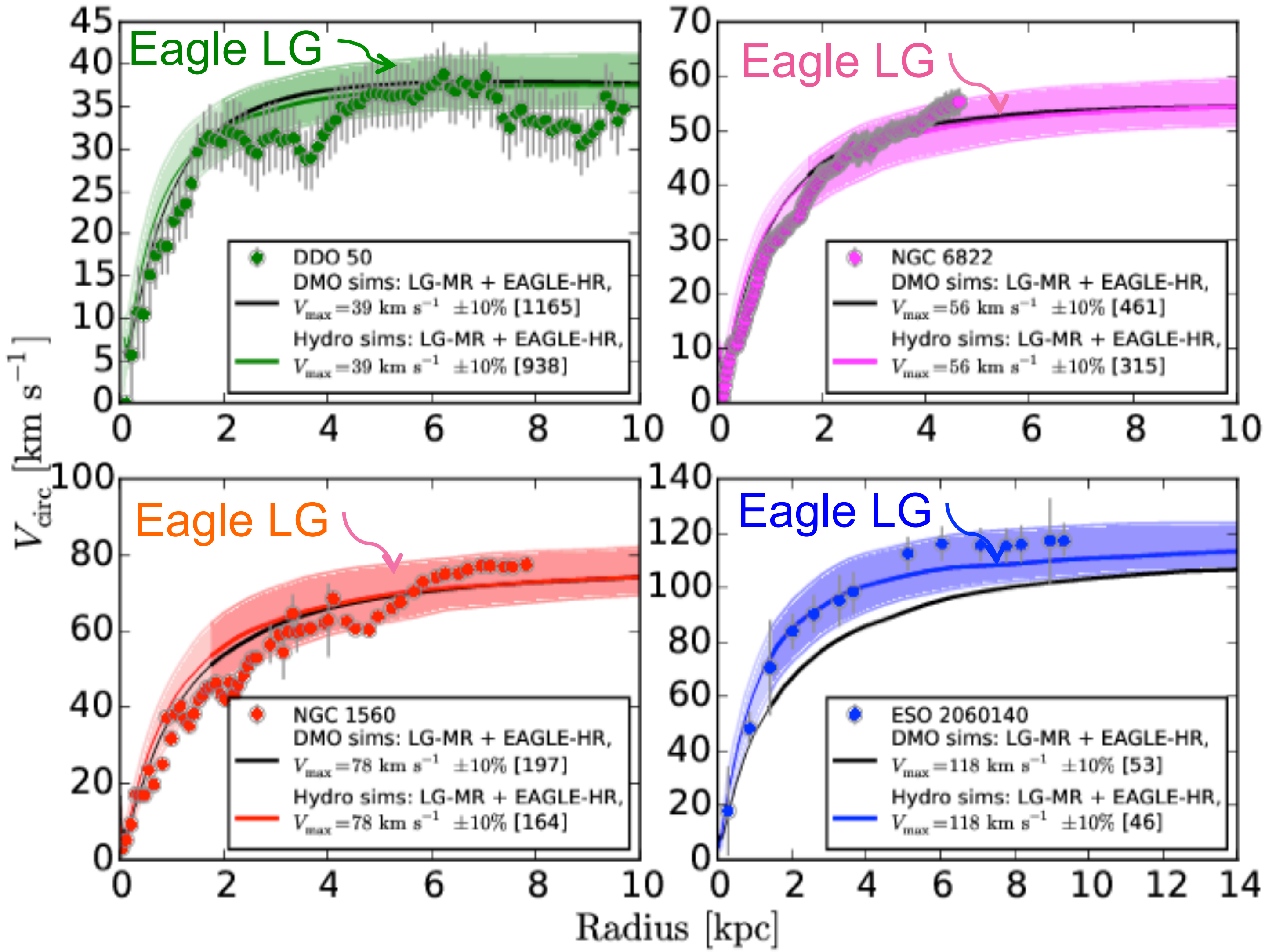
(i) Rotation curves

(ii) 2D velocity fields

The diversity of gal rotation curves

Four rotation curves that are well fit by ΛCDM

(from dwarfs to ~L_*)

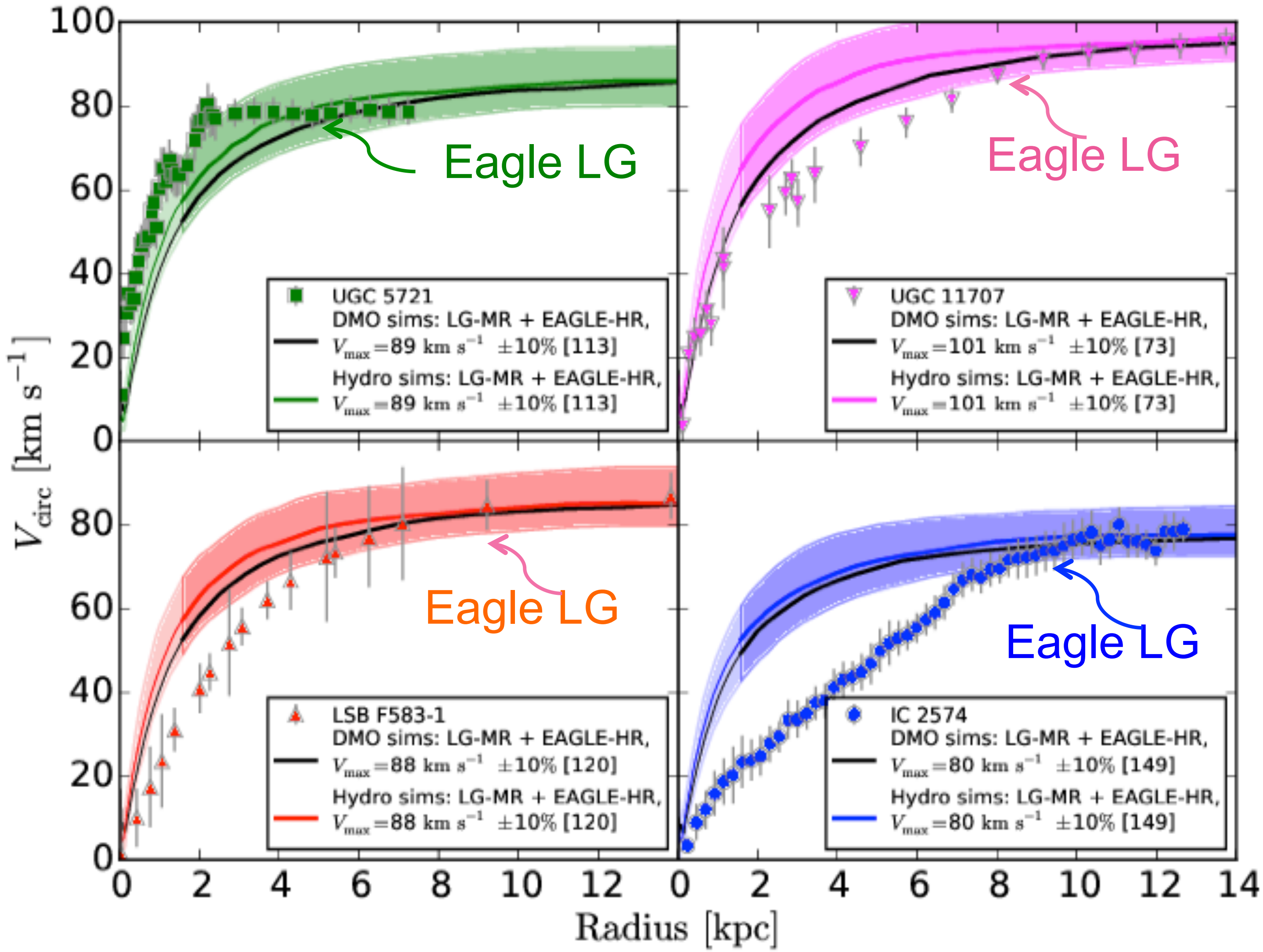


Oman, Navarro, Frenk et al. ‘15

University of Durham

The diversity of gal rotation curves

Four rotation curves that are NOT well fit by ΛCDM

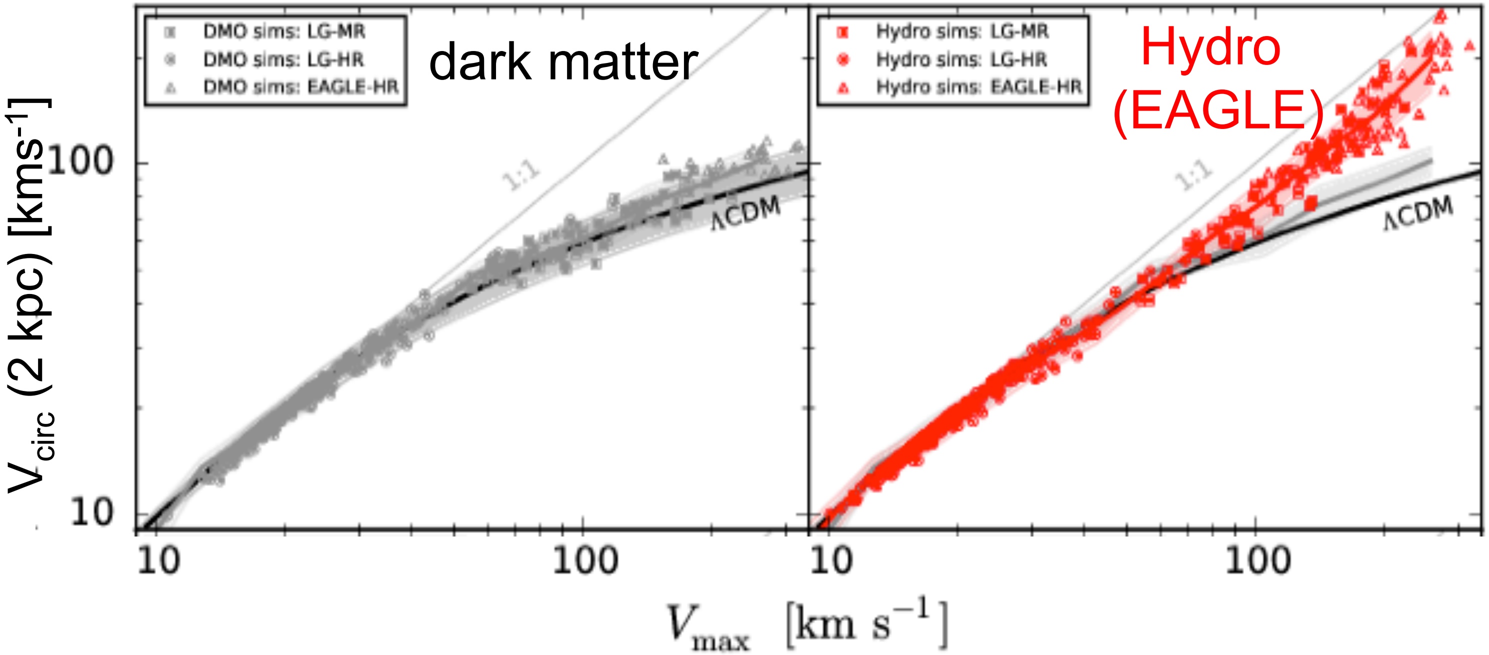


(from dwarfs to ~L_*)

Oman et al. '15

Institute for Computational Cosmology

The diversity of gal rotation curves



Oman et al. '15

The diversity of gal rotation curves

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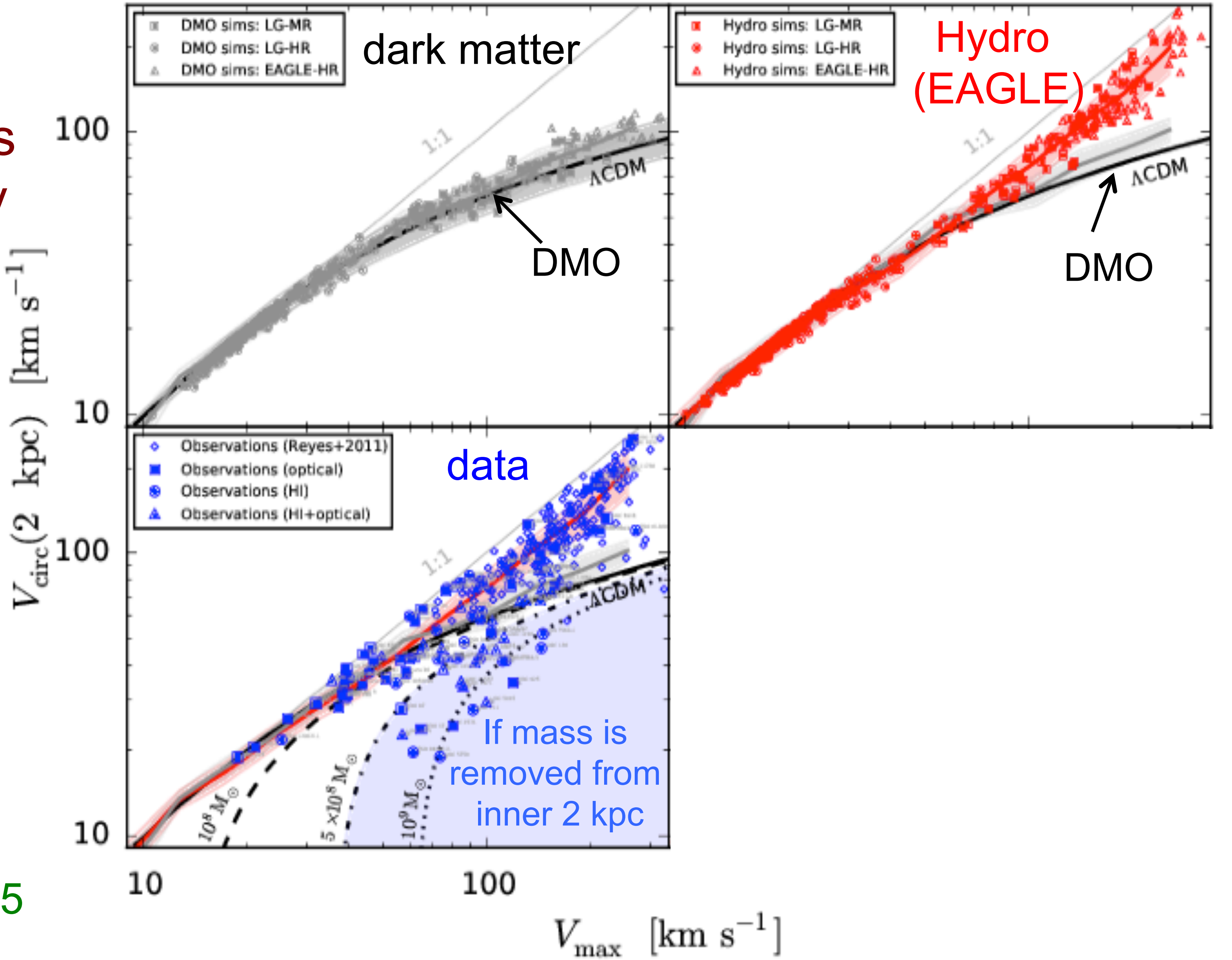
Most galaxies are well fit by EAGLE; others not fit by any simulation



Oman et al. '15

Are there other baryon effects that could make cores but are not present in Eagle?

Mon. Not. R. Astron. Soc. **283**, L72–L78 (1996)

The cores of dwarf galaxy haloes

Julio F. Navarro,[1,2]★ Vincent R. Eke[2] and Carlos S. Frenk[2]

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ABSTRACT

We use *N*-body simulations to examine the effects of mass outflows on the density profiles of cold dark matter (CDM) haloes surrounding dwarf galaxies. In particular, we investigate the consequences of supernova-driven winds that expel a large fraction of the baryonic component from a dwarf galaxy disc after a vigorous episode of star formation. We show that this sudden loss of mass leads to the formation of a core in the dark matter density profile, although the original halo is modelled by a coreless (Hernquist) profile. The core radius thus created is a sensitive function of the mass and radius of the baryonic disc being blown up. The loss of a disc with mass and size consistent with primordial nucleosynthesis constraints and angular momentum considerations imprints a core radius that is only a small fraction of the original scalelength of the halo. These small perturbations are, however, enough to reconcile the rotation curves of dwarf irregulars with the density profiles of haloes formed in the standard CDM scenario.

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Baryon effects in the MW satellites

Let gas cool and condense to the galactic centre

→ gas self-gravitating
→ star formation/burst

Rapid ejection of gas during starburst → a core in the halo dark matter density profile

Navarro, Eke, Frenk '96

Governato et al. '12
Pontzen & Governato '12
Brooks et al. '12

Navarro, Eke, Frenk '96

The cores of dwarf galaxy haloes L75

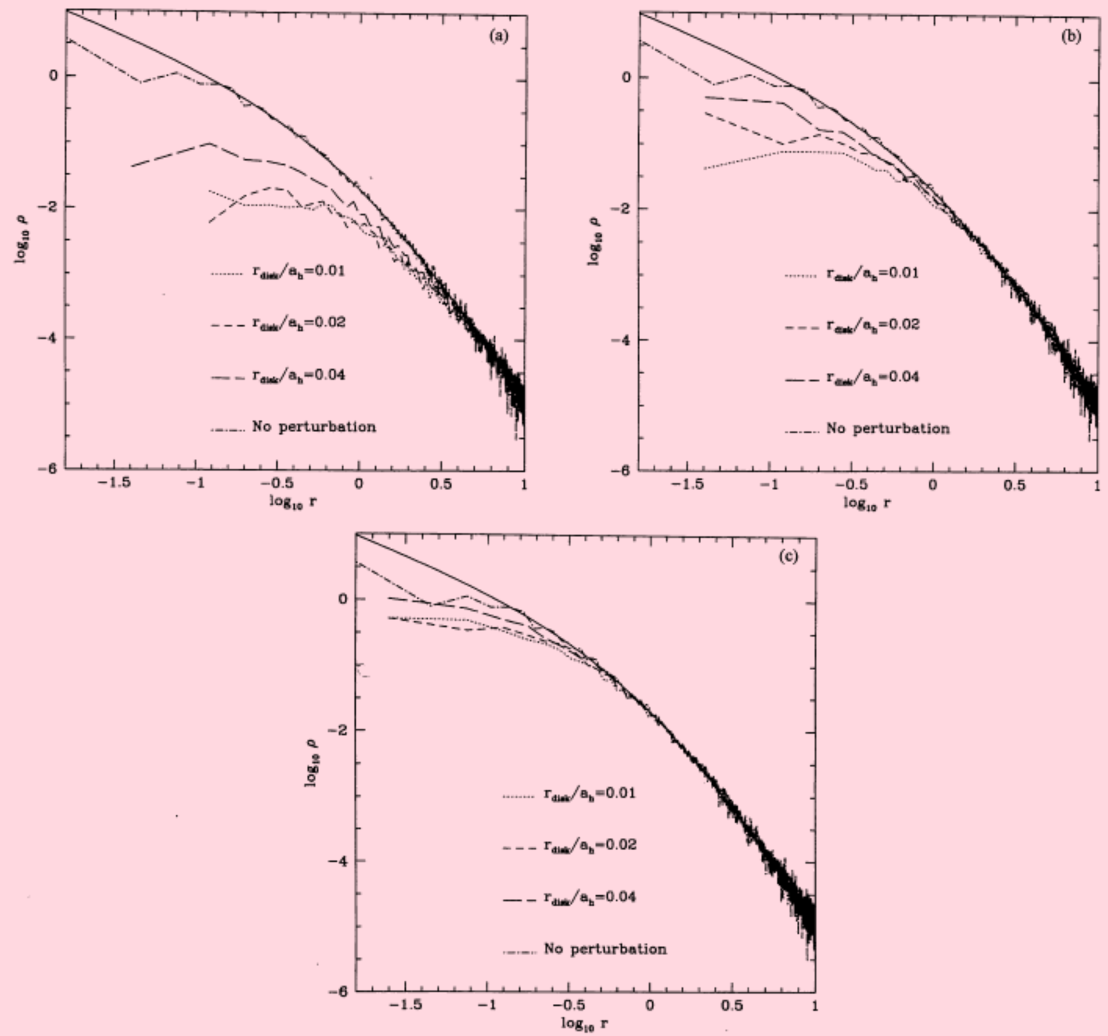

Figure 3. Equilibrium density profiles of haloes after removal of the disc. The solid line is the original Hernquist profile, common to all cases. The dot–dashed line is the equilibrium profile of the 10 000-particle realization of the Hernquist model run in isolation at $t = 200$. (a) $M_{disc} = 0.2$. (b) $M_{disc} = 0.1$. (c) $M_{disc} = 0.05$.

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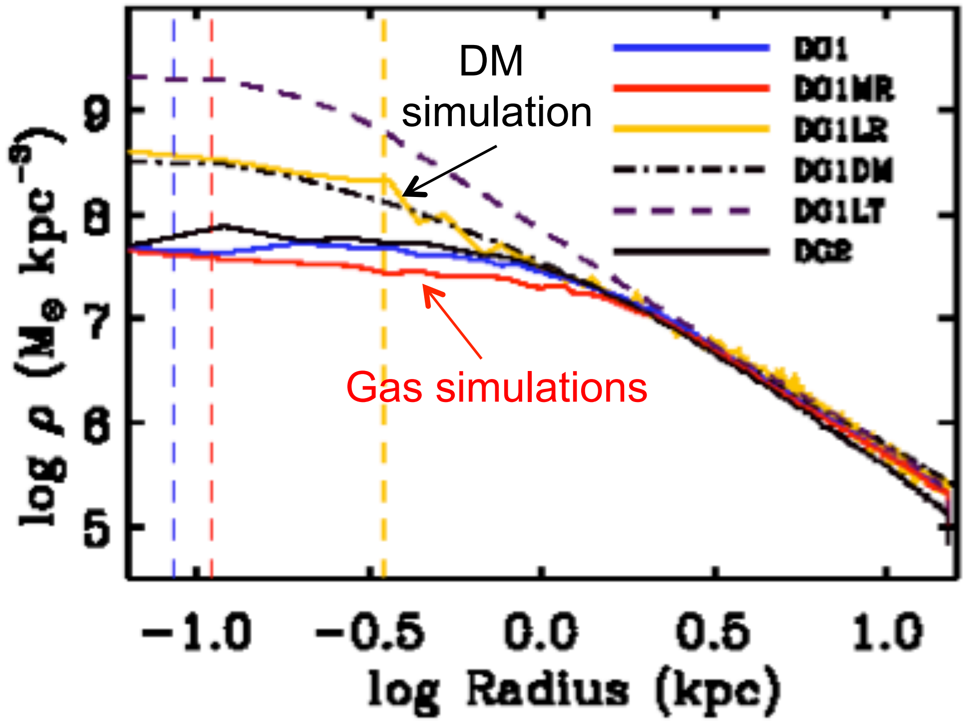
Cores in dwarf galaxy simulations

Governato et al. assume high density threshold for star formation.

EAGLE does not

→ High threshold allows large gas mass to accummulate in centre

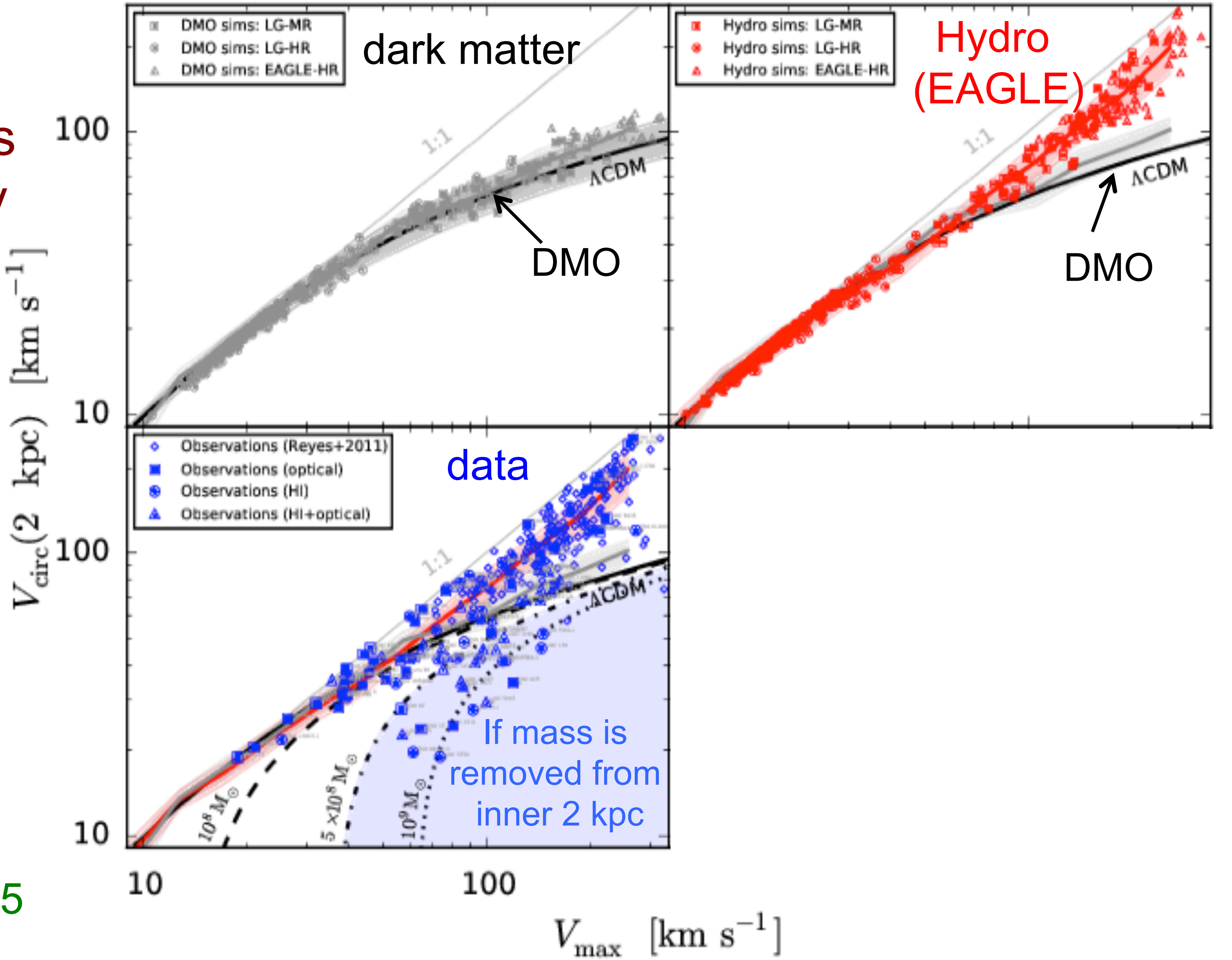
→ Sudden repeated removal of gas transfers binding energy



Governato et al. ’10
Pontzen et al. 2011

The diversity of gal rotation curves

Most galaxies are well fit by EAGLE; others not fit by any simulation



Oman et al. '15

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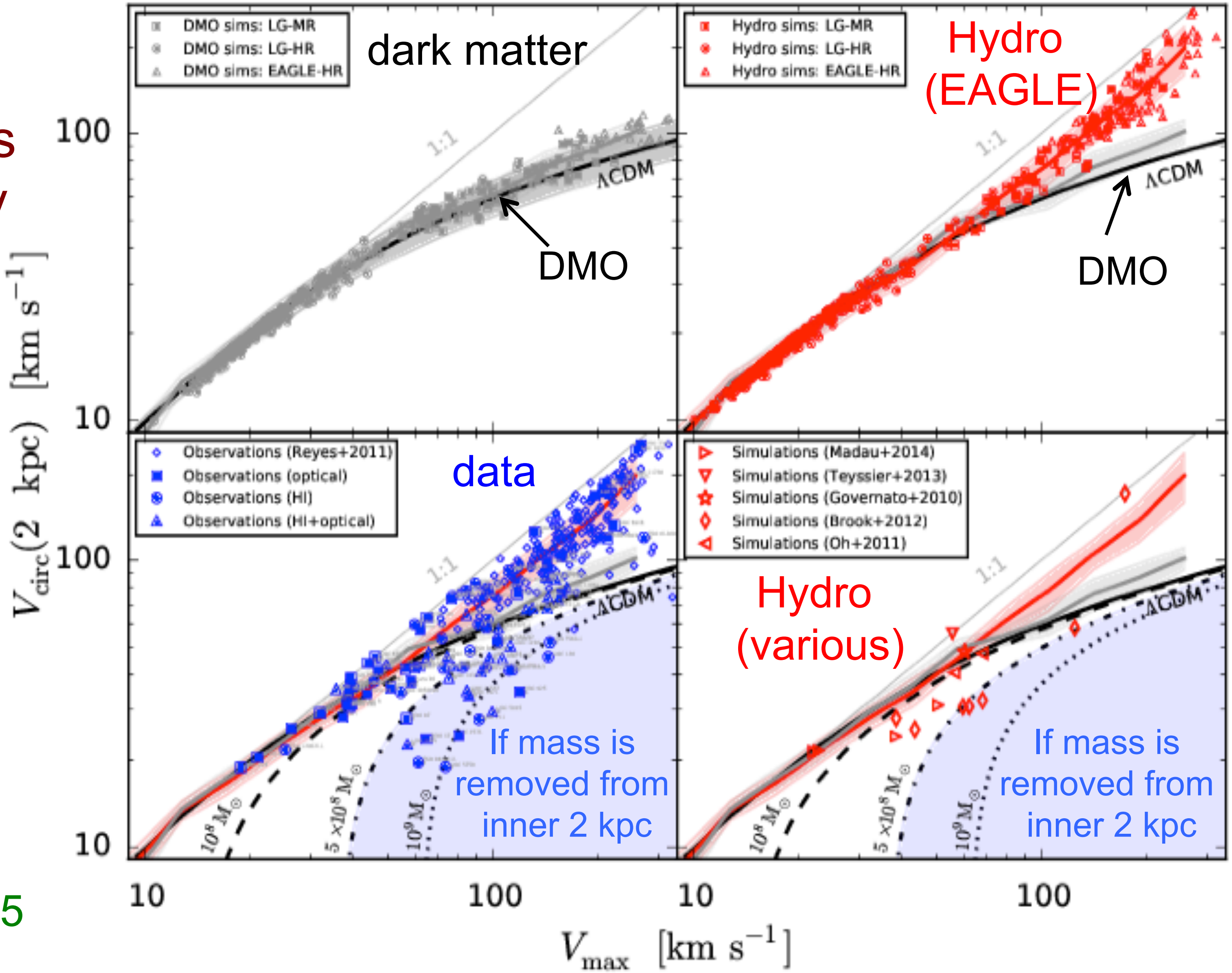

Oman et al. '15

Cores or cusps in dwarf gals?

- Some dwarfs have rotation curves that agree well with EAGLE
- Others have inner mass deficits compared to ΛCDM expectation
- In many cases, inner deficit much larger than seen in simulations that make cores

EITHER (i) dark matter more complex than in any current model

OR (ii) current simulations fail to reproduce effects of baryons on inner regions of dwarfs

AND/OR (iii) the mass profiles of "inner mass deficit" galaxies inferred from kinematic data are incorrect.

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Four problems for CDM on small scales?

1. The “missing satellites” problem
2. The “too-big-to-fail” problem
3. The “core-cusp” problem
4. The “satellite disk” problem

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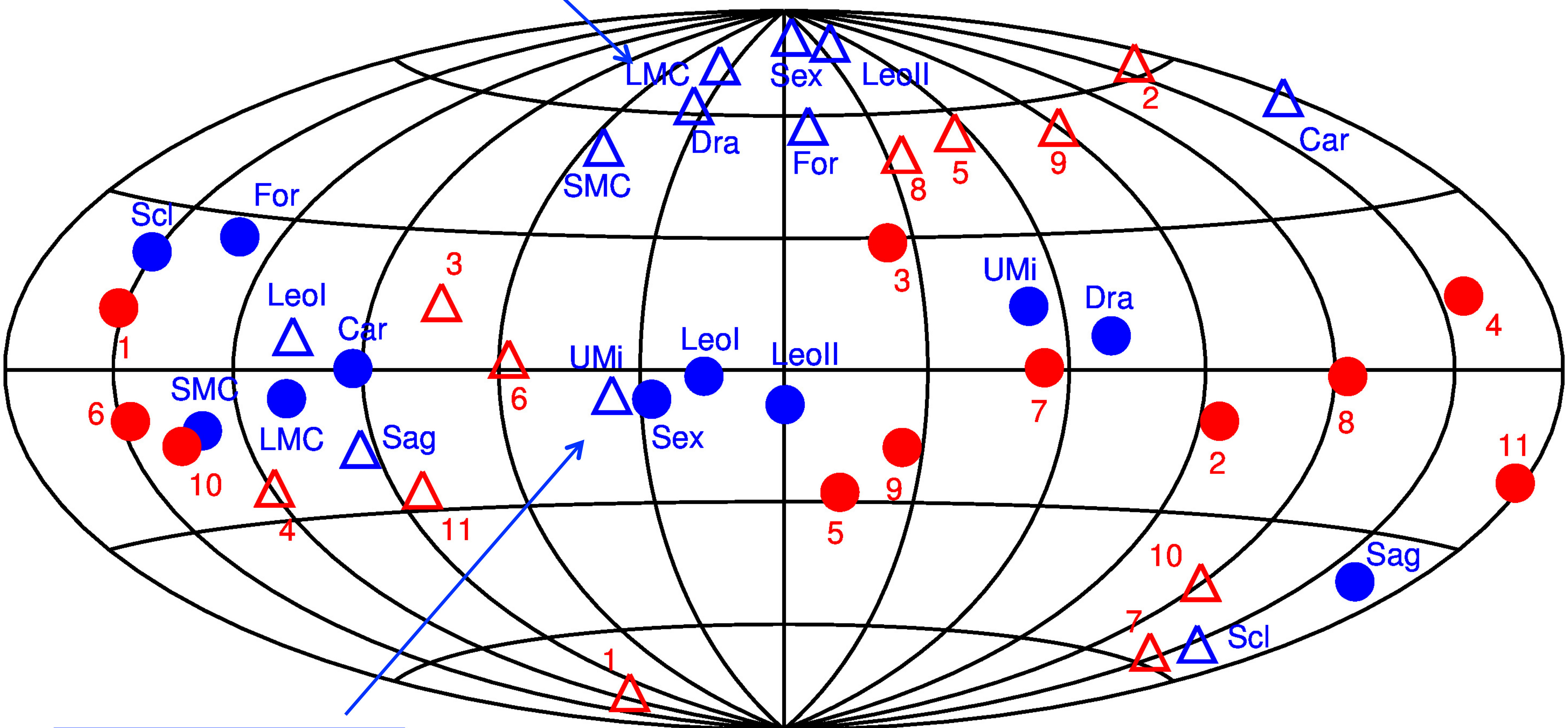
The "satellite disk" problem

Direction of ang. mom.

Milky Way



MW satellites

Lynden-Bell '76

Sawala et al '15

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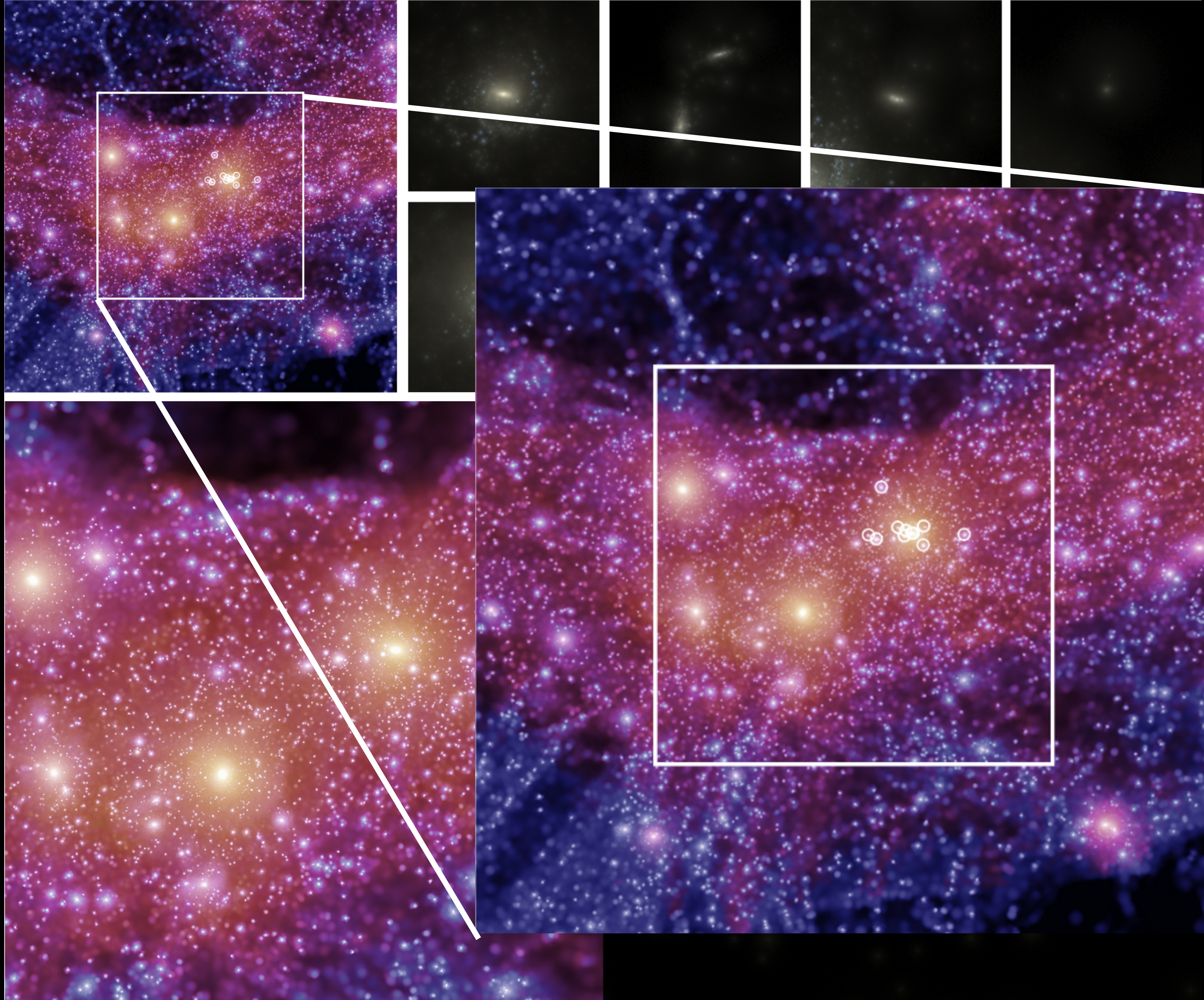
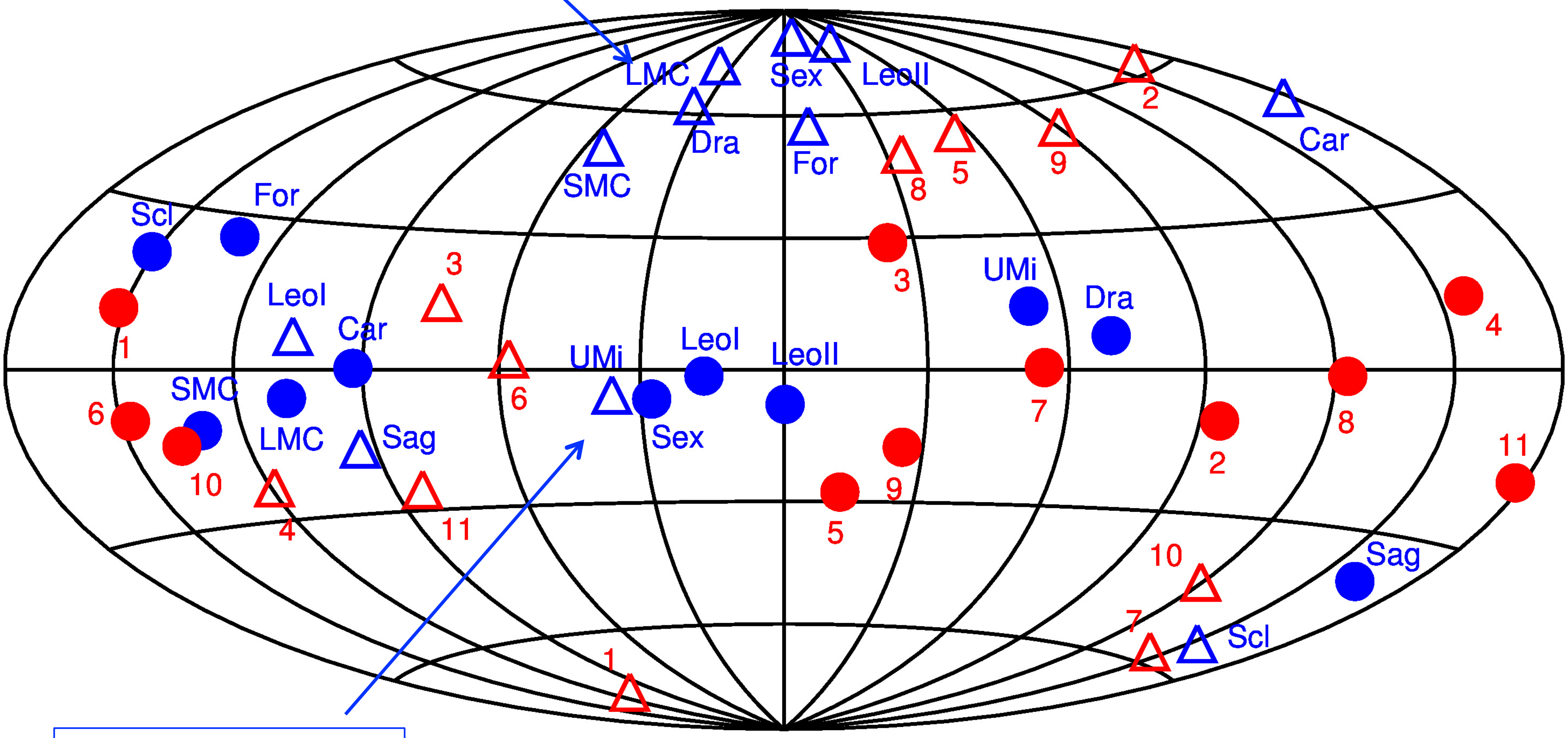


The "satellite disk" problem

University of Durham

Direction of ang. mom.

Milky Way



MW satellites

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Sawala et al '15

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The curious case of a thin, "rotating" plane of sats in M31

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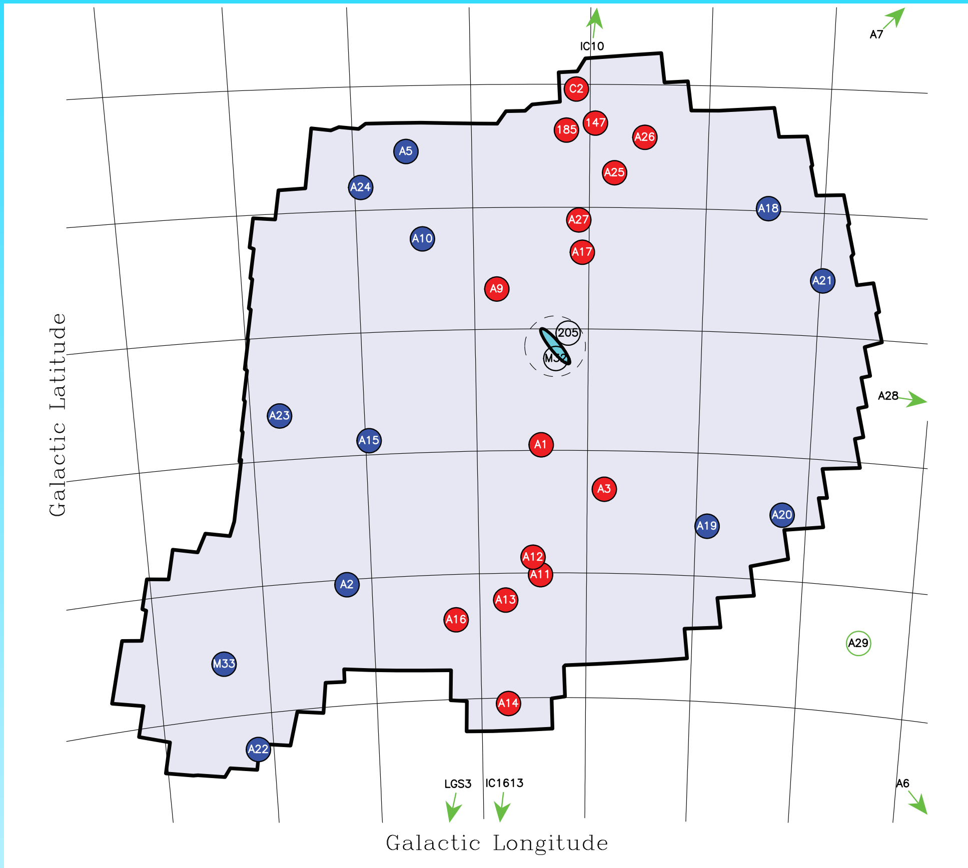
Ibata et al '13 found a plane of 15 satellites in Andromeda (out of 27) of which 13 have the same sense of rotation

They claim a 4.3σ detection

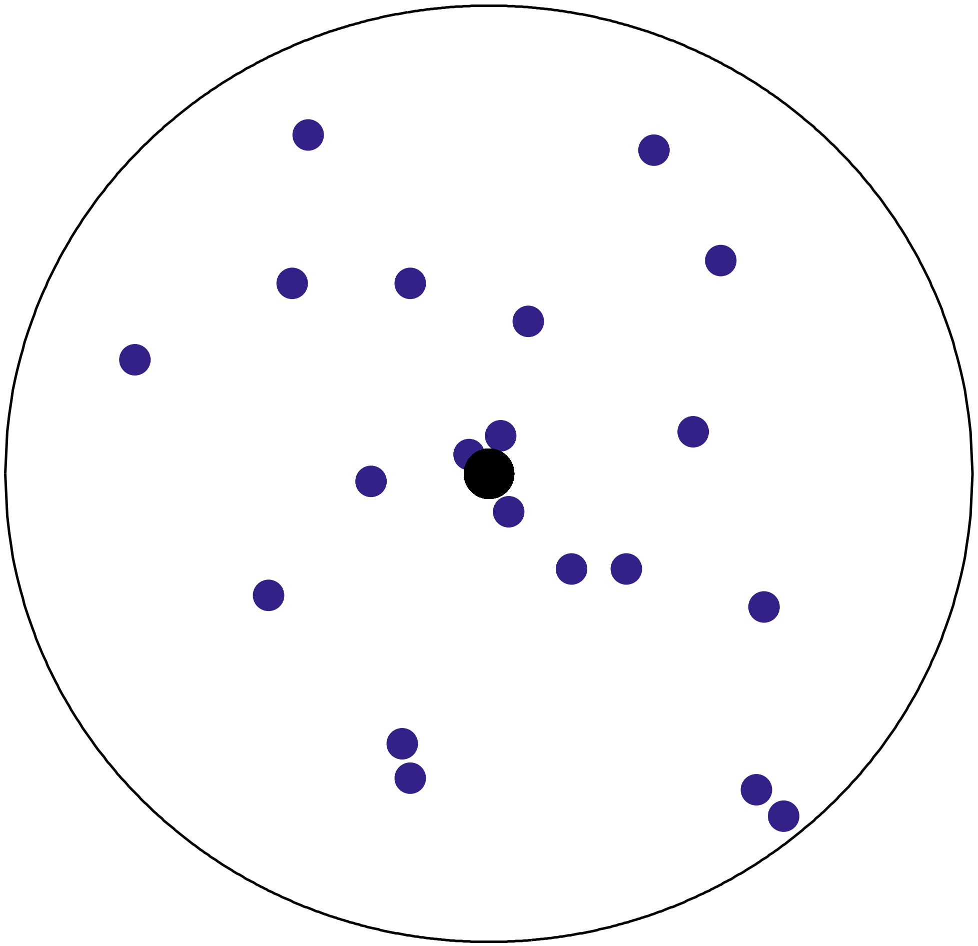
"We find that 0.04% of host galaxies [in Millennium II] display satellite alignments that are at least as extreme as the observations, when we consider their extent, thickness, and number of members rotating in the same sense."



Ibata et al '13

University of Durham

Finding disks of satellites



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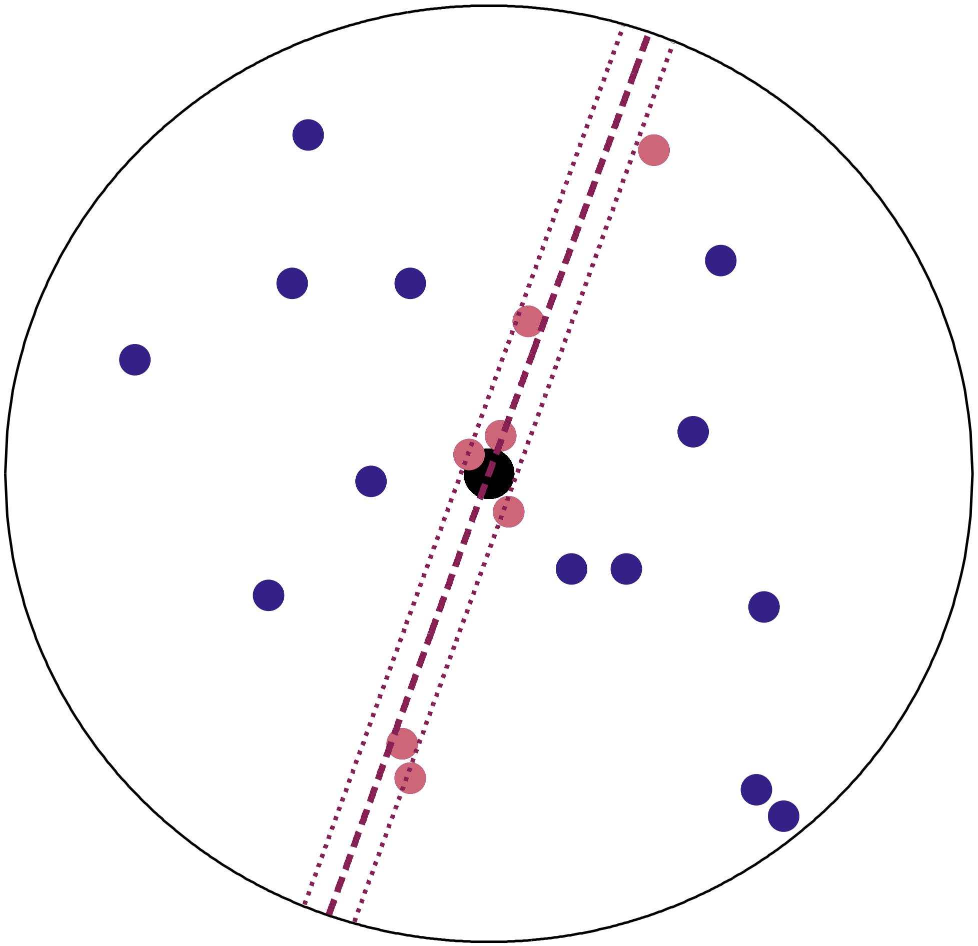
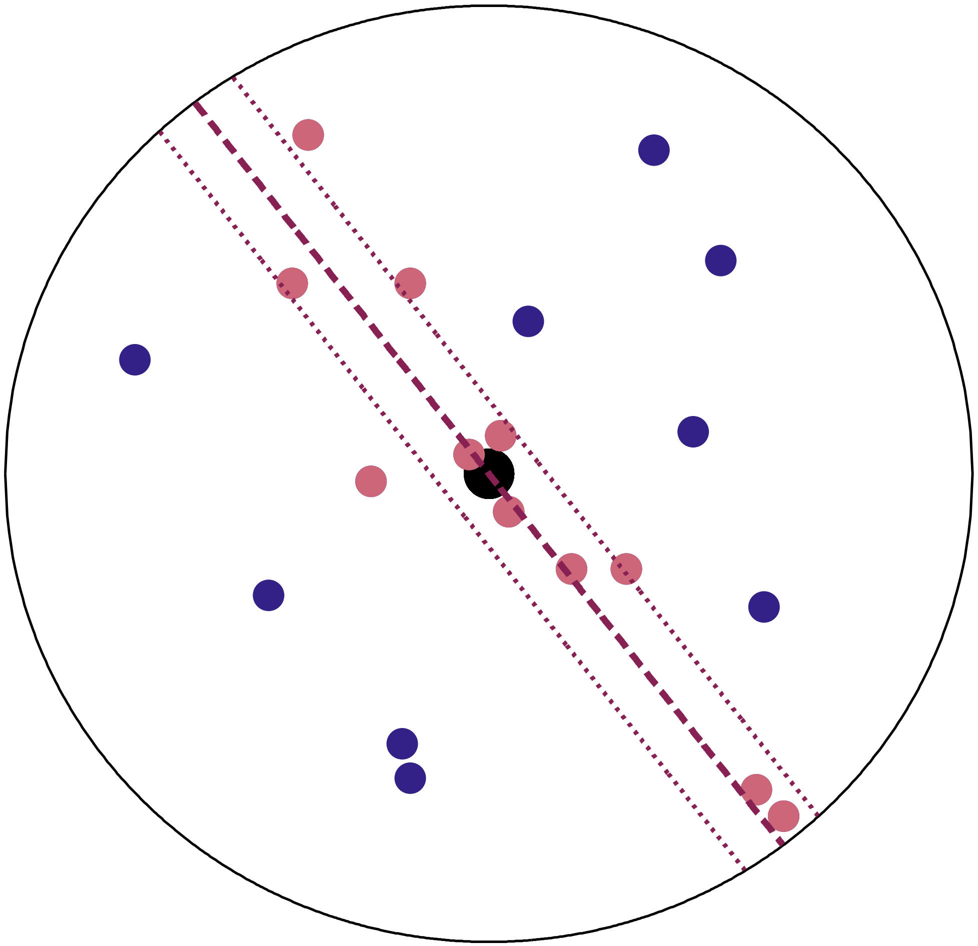
Finding disks of satellites

Plane 1: $N_{sat} = 7$, $\mathcal{P} = 410$



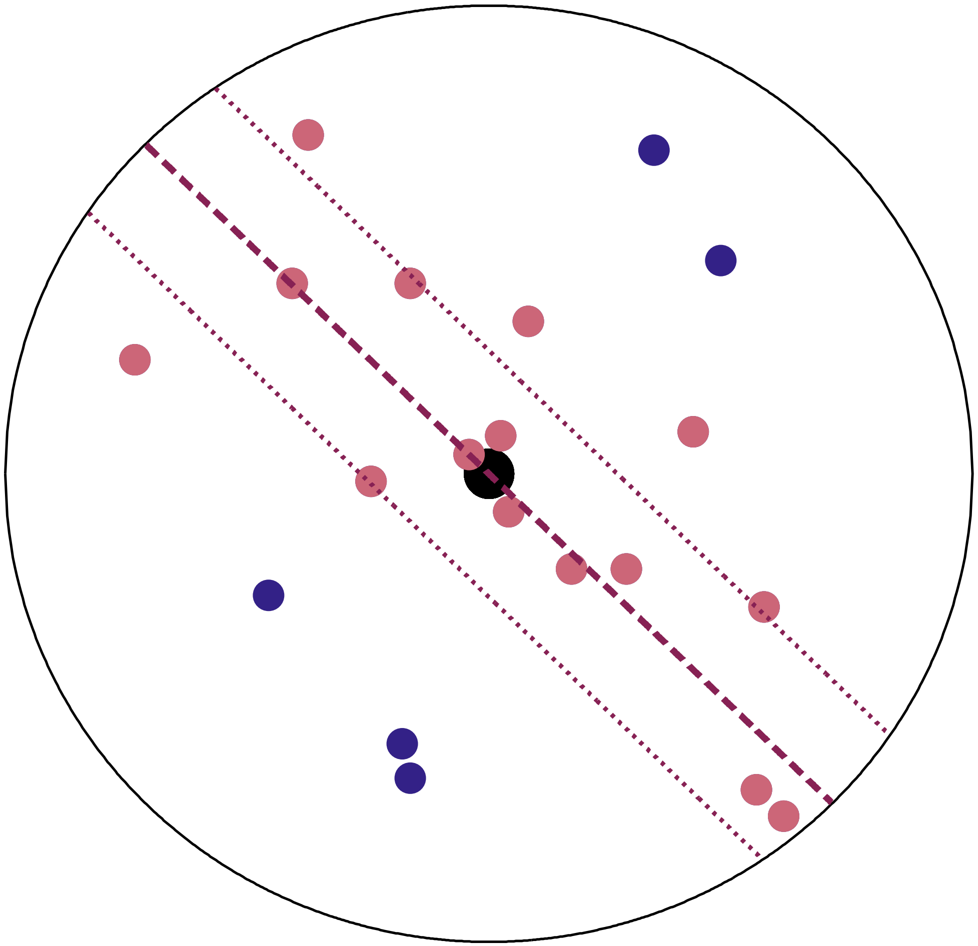
Finding disks of satellites

Plane 2: $N_{sat} = 11,\ \mathcal{P} = 660$



Finding disks of satellites

Plane 3: $N_{sat} = 15$, $\mathcal{P} = 450$



The "satellite disk" problem

Prominence of a plane = $\frac{1}{\text{Probability of finding plane in random distr}}$

Prominence of plane thiner than $r_\perp$ having N_{sat} galaxies

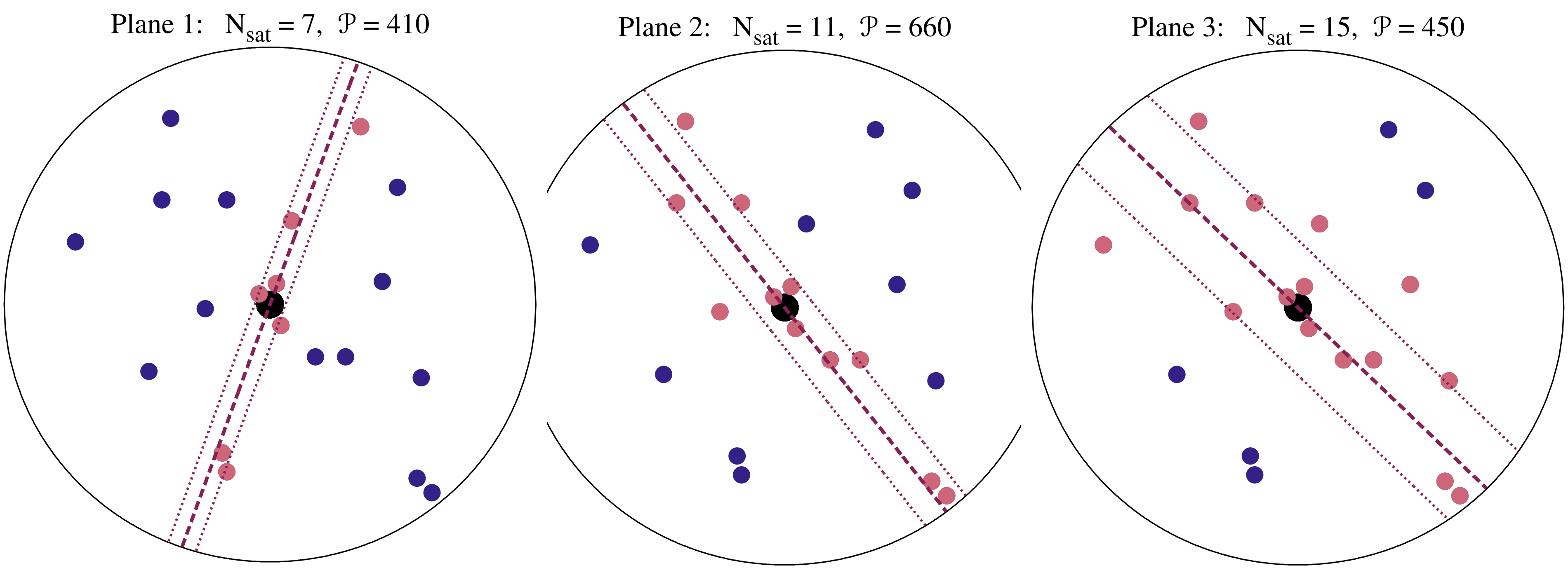
$$\mathcal{P}^{\mathrm{plane\ i}}_{\mathrm{spatial}} = \frac{1}{p\left(\leqslant r_{\perp;\, i} \mid N_{\mathrm{sat;\ i}}\right)}$$

Prominence of plane of N_{sat} gals, N same sense of rotation

$$\mathcal{P}^{\mathrm{plane\ i}}_{\mathrm{2D-kin}} = \frac{1}{p\left(\geqslant N_{\mathrm{s.s.r.;\ i}} \mid N_{\mathrm{sat;\ i}}\right)}$$

$$\mathcal{P}^{\mathrm{rarest}}_{\mathrm{spatial}} = \max_{\mathrm{all\ planes\ i}} \left[\mathcal{P}^{\mathrm{plane\ i}}_{\mathrm{spatial}} \right] ;$$

Finding disks of satellites



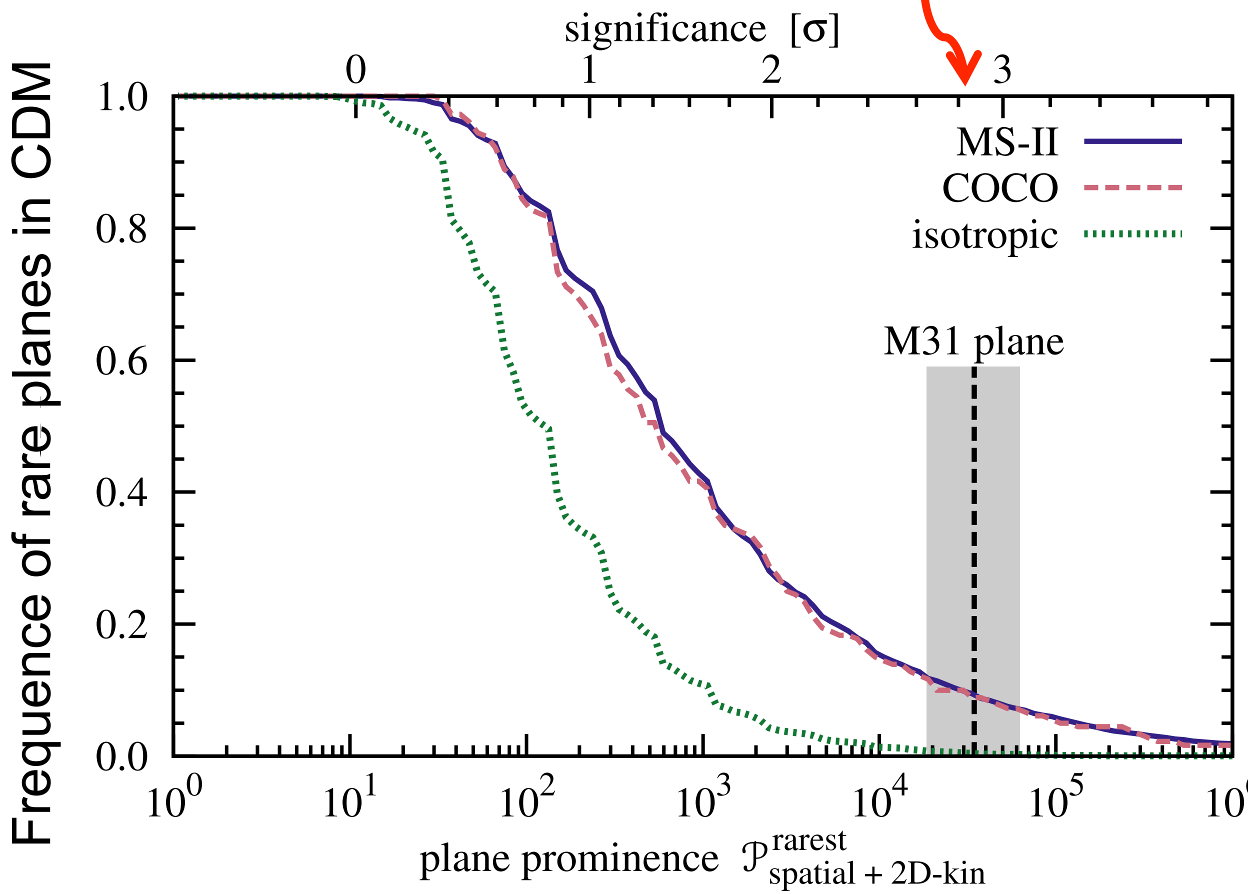
Cautun, Bose, Frenk et al ‘15

The significance of Ibata's plane

University of Durham

- Significance of Ibata's plane is reduced by x100 when trials factor is included
- 8.8% of halos in ΛCDM simulation have even more prominent disks than Ibata's

Including "trials factor"



Cautun et al '15

1 in 30,000 chance of finding a plane of 15 sats (out of 27) as thin as the one found by Ibata et al., with at least 13 having the same sense of rotation

Is there a "satellite disks problem" in CDM?

No, when statistics are properly calculated

Satellite planes are v. common in ΛCDM: 5 & 9% of halos have even more prominent planes than Milky Way and Andromeda

Conclusions

- ΛCDM: great success on scales > 1Mpc: CMB, LSS, gal evolution

Four “problems” on small scales:

1. Abundance of sats
2. Too-big-to-fail
3. Core-cusp
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Four “problems” on small scales:

1. Core-cusp: Not a problem? (Baryon effects?)
2. Abundance of sats: CDM OK; WDM OK if m_{WDM}> 3 KeV
3. Too-big-to-fail: CDM, WDM OK
4. Disk of satellites: CDM, WDM OK